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## Citation

Tomsick, John A., Eric V. Gotthelf, Farid Rahoui, Roberto J. Assef, Franz E. Bauer, Arash Bodaghee, Steven E. Boggs, et al. 2014. "NuSTAR J163433-4738.7: A Fast X-Ray Transient in the Galactic Plane." *The Astrophysical Journal* 785 (1) (March 19): 4. doi:10.1088/0004-637x/785/1/4.

## Published Version

doi:10.1088/0004-637X/785/1/4

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## NUSTAR J163433–4738.7: A FAST X-RAY TRANSIENT IN THE GALACTIC PLANE

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Accepted by ApJ

### ABSTRACT

During hard X-ray observations of the Norma spiral arm region by the *Nuclear Spectroscopic Telescope Array* (*NuSTAR*) in 2013 February, a new transient source, NuSTAR J163433–4738.7, was detected at a significance level of  $8\text{-}\sigma$  in the 3–10 keV bandpass. The source is consistent with having a constant *NuSTAR* count rate over a period of 40 ks and is also detected simultaneously by *Swift* at lower significance. The source is not significantly detected by *NuSTAR*, *Swift*, or *Chandra* in the days before or weeks after the discovery of the transient, indicating that the strong X-ray activity lasted for between  $\sim 0.5$  and 1.5 days. Near-IR imaging observations were carried out before and after the X-ray activity, but we are not able to identify the counterpart. The combined *NuSTAR* and *Swift* energy spectrum is consistent with a power-law with a photon index of  $\Gamma = 4.1_{-1.0}^{+1.5}$  (90% confidence errors), a blackbody with  $kT = 1.2 \pm 0.3$  keV, or a bremsstrahlung model with  $kT = 3.0_{-1.2}^{+2.1}$  keV. The reduced- $\chi^2$  values for the three models are not significantly different, ranging from 1.23 to 1.44 for 8 degrees of freedom. The spectrum is strongly absorbed with  $N_{\text{H}} = (2.8_{-1.4}^{+2.3}) \times 10^{23}$  cm<sup>-2</sup>,  $(9_{-7}^{+15}) \times 10^{22}$  cm<sup>-2</sup>, and  $(1.7_{-0.9}^{+1.7}) \times 10^{23}$  cm<sup>-2</sup>, for the power-law, blackbody, and bremsstrahlung models, respectively. Although the high column density could be due to material local to the source, it is consistent with absorption from interstellar material along the line of sight at a distance of 11 kpc, which would indicate an X-ray luminosity  $> 10^{34}$  erg s<sup>-1</sup>. Although we do not reach a definitive determination of the nature of NuSTAR J163433–4738.7, we suggest that it may be an unusually bright active binary or a magnetar.

*Subject headings:* stars: variables: general — X-rays: stars — surveys — Galaxy: stellar content — X-rays: individual (NuSTAR J163433–4738.7)

### 1. INTRODUCTION

Hard X-ray surveys of the Galaxy provide an opportunity to discover populations of extreme sources. The promise of such surveys has been partially realized with the *International Gamma-Ray Astrophysics Laboratory* (*INTEGRAL*, Winkler et al. 2003), which carried out a 20–100 keV survey

of the entire Galactic Plane and has discovered hundreds of new sources (Bird et al. 2010; Krivonos et al. 2012), including new types of High Mass X-ray Binaries (HMXBs), pulsar wind nebulae (PWNe), and magnetic Cataclysmic Variables (CVs). The *Nuclear Spectroscopic Telescope Array* (*NuSTAR*, Harrison et al. 2013), which launched in 2012 June and covers the 3–79 keV bandpass, is the first focusing hard X-ray telescope in orbit. While it has a much smaller field of view than *INTEGRAL*, it has much lower background and greatly improved sensitivity. Thus, one of *NuSTAR*'s science goals is to extend our view deeper into the Galactic plane to look for hidden hard X-ray populations.

During its first year of operation, *NuSTAR* has initiated surveys of  $\sim 1$  deg<sup>2</sup> areas both in the Galactic Center (e.g., Mori et al. 2013; Nynka et al. 2013) and in a region which samples the spiral arm population in order to probe potentially different environments where X-ray binaries are found. A spiral arm region centered on Galactic coordinates of  $l = 337.5^\circ$  and  $b = 0^\circ$  was chosen for having the highest known density of OB star associations (Russeil 2003) and HMXBs (Bodaghee et al. 2007, 2012). This is part of the “Norma” spiral arm region, which was identified early in the *INTEGRAL* mission as having an unusually high density of hard X-ray sources (Tomsick et al. 2004; Dean et al. 2005; Lutovinov et al. 2005). The combination of active star formation and evidence that compact objects have already formed suggests that a survey by *NuSTAR* may uncover compact objects associated with populations of massive stars such as magnetars or faint HMXBs that are early in their evolutionary process and may have neutron star or black hole accretors.

The full surveys of the Galactic Center and the Norma region will be carried out over a period of  $\sim 2$  years; here we

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report on the discovery of a transient source made during the first part of the survey. In the following, §2 describes observations made with *NuSTAR*, *Swift*, and *Chandra*, as well as the procedures we used to reduce the data. The results are presented in §3, and we discuss possibilities for the nature of the transient in §4.

## 2. OBSERVATIONS AND DATA REDUCTION

As the first part of the *NuSTAR* survey of the Norma region, nine  $\sim 20$  ks *NuSTAR* observations were performed between UT 2013 February 20 and February 24. Each  $13'$ -by- $13'$  field of view was partially overlapping with adjacent pointings, and the entire region covered was  $\sim 0.2$  deg<sup>2</sup>. The results from all nine pointings will be reported in Bodaghee et al. (submitted). Here we focus on the observations that covered a new transient, NuSTAR J163433–4738.7. These observations are listed in Table 1, including two that were obtained during the survey and a follow-up *NuSTAR* observation on 2013 March 23 that was coordinated with *Chandra*.

In addition,  $\sim 2$  ks *Swift* X-ray Telescope (XRT) observations of the region were carried out during 2013 February 21–24, and four of these observations covered NuSTAR J163433–4738.7. Table 1 lists these along with three other *Swift* observations that were acquired after our survey as part of another observing program. We also analyzed archival data covering the source, including two *Chandra* observations (ObsIDs 12529 and 12532 with exposure times of 19.0 ks and 19.5 ks, respectively) that were acquired in 2011 as part of a survey of the same region being covered by *NuSTAR* (Fornasini et al., submitted).

We reduced the *NuSTAR* and *Swift* data using HEASOFT v6.14 and the latest version of the Calibration Database (CALDB) files as of 2013 August 30. We produced cleaned event lists for the *NuSTAR* Focal Plane Modules (FPMA and FPMB) using *nupipeline* and for the *Swift*/XRT using *xrtpipeline*, and further analysis of the event lists is described below. For *Chandra*, we processed the Advanced CCD Imaging Spectrometer (ACIS, Garmire et al. 2003) data with the *Chandra* Interactive Analysis of Observations (CIAO) software, using *chandra\_repro* to make event lists.

We obtained near-IR observations covering the NuSTAR J163433–4738.7 error region. This includes *J*, *H* and *K<sub>s</sub>* observations performed on 2011 July 19 with CTIO/NEWFIRM in the framework of near-IR mapping of the *Chandra* survey field. A detailed description of the data and their reduction can be found in Rahoui et al. (in prep). They were reduced with the dedicated IRAF package *NFEXTERN* following the standard procedure – tailored for wide-field mosaics – which consists of bad pixel removal, dark subtraction, linearity correction, flatfielding and median sky subtraction. The resulting images were then flux-calibrated through relative photometry with the 2MASS catalogue. We also obtained *K<sub>s</sub>*-band imaging with the Ohio-State Infra-Red Imager/Spectrometer (OSIRIS) at the Southern Astrophysical Research (SOAR) 4.1m Telescope. We used the f/7 camera, providing a  $80''$  field of view, centered at the nominal coordinates of NuSTAR J163433–4738.7. We obtained  $36 \times 60$  s dithered exposures of the field on 2013 April 3 under good conditions with seeing of  $0.''7$ , and observed it for  $9 \times 60$  s again on 2013 April 5, also under good conditions but with somewhat worse seeing of  $1.''0$ . Reductions were done using the *XDIMSUM* IRAF package and photometric calibration was obtained by comparing to 2MASS All-Sky Point Source

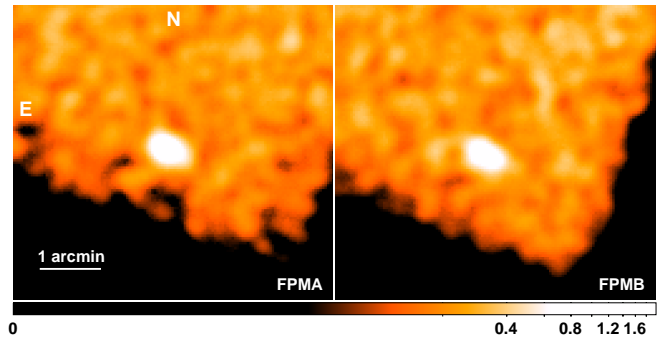


FIG. 1.— *NuSTAR* discovery images in the 3–20 keV energy band for NuSTAR J163433–4738.7. The source was detected in both of the *NuSTAR* Focal Plane Modules (FPMA and FPMB) in a 22.6 ks exposure taken on 2013 Feb. 23–24. The images have been rebinned so that the pixel size is  $2.''5$  and smoothed with a 6-pixel Gaussian. The scale on the bottom of the figure is in counts per pixel, and a logarithmic scaling is used. The apparent elongation of the source is due to the distorted PSF shape at large off-axis angles.

Catalog sources in the field.

## 3. RESULTS

The new transient was discovered from an inspection of the image from the 22.6 ks *NuSTAR* observation that took place starting on 2013 February 23, 14.52 h. As shown in Figure 1, the source was detected in FPMA and FPMB. To determine the significance of the detection, we extracted 3–10 keV counts from a  $30''$ -radius circle centered on the approximate position of the source. We determined the background level using a nearby source-free circular region with a radius of  $90''$ . After background subtraction, we obtained  $97 \pm 14$  and  $99 \pm 20$  counts in FPMA and FPMB, respectively. The combined significance is  $8.0\text{-}\sigma$ , confirming the detection.

To determine the source position, we extracted all the events from a  $60''$ -by- $60''$  square region and made histograms of the counts, binning in the R.A. and Decl. directions. We performed  $\chi^2$  fitting of the histograms with a model consisting of a constant (accounting for the flat background) and a Gaussian for the source. It should be noted that this is an approximation since the *NuSTAR* point spread function (PSF) is non-Gaussian. We determined the centroids separately for FPMA and FPMB, and they are consistent with each other. The weighted average of the two centroids is R.A. =  $16^{\text{h}}34^{\text{m}}33.^{\text{s}}42$ , Decl. =  $-47^{\circ}38'41.''9$  (J2000.0) with  $3\text{-}\sigma$  statistical uncertainties of  $6.''3$  and  $4.''9$  in R.A. and Decl., respectively. After considering that the systematic pointing uncertainty for *NuSTAR* is  $\sim 8''$  (Harrison et al. 2013), the error region can be approximated with a  $10''$ -radius circle.

We searched on-line catalogs (e.g., SIMBAD<sup>16</sup>) for X-ray sources consistent with the position of NuSTAR J163433–4738.7, but we did not find any likely candidates. The catalog from the 2011 *Chandra* survey (Fornasini et al., submitted) does not have any sources in the NuSTAR J163433–4738.7 error region. Based on this and the analysis described below (including a re-analysis of the 2011 *Chandra* observations), we conclude that NuSTAR J163433–4738.7 is a previously undetected source and that it is very likely to be a transient given the sensitivity of the *Chandra* observations.

For the *NuSTAR* and *Swift* observations listed in Table 1, we estimated count rates or upper limits for NuSTAR J163433–

<sup>16</sup> <http://simbad.u-strasbg.fr/simbad/>

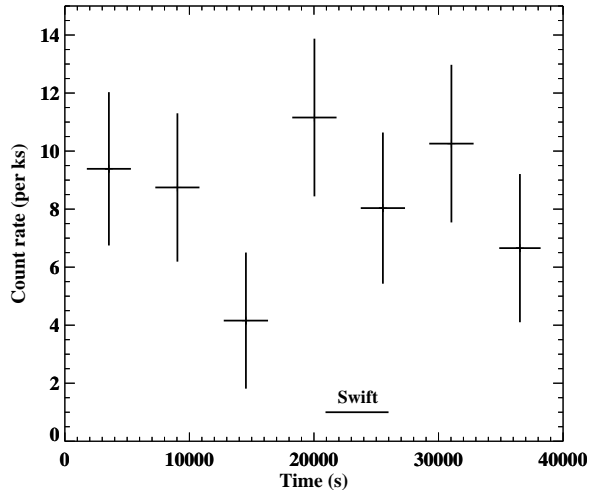


FIG. 2.— The 3–10 keV *NuSTAR* light curves (FPMA and FPMB after background subtraction) for NuSTAR J163433–4738.7 during ObsID 40014007001. There is one data point per satellite orbit. The errors on the data points are at  $1\text{-}\sigma$  confidence. The time of a *Swift* observation is indicated. Zero on the time axis corresponds to MJD 56,346.60000.

4738.7 using  $30''$ -radius source regions centered at the source position derived from ObsID 40014007001. The XRT 90% encircled energy radius is approximately  $20''$ , but we use  $30''$  to account for the uncertainty in the source position as NuSTAR J163433–4738.7 is not bright enough in any of the *Swift* observations to improve the measurement of the source position. As described above, we used a larger, circular region that does not contain any detected point sources for background. Selecting the background regions for *NuSTAR* requires some care because of scattered light from a nearby bright source (4U 1630–47).

The 3–10 keV count rates or limits obtained for all observations are given in Table 1. For *NuSTAR*, the background rates are high enough to use Gaussian statistics, and we use the Poisson limits tabulated in Gehrels (1986) for *Swift* and *Chandra*. In all cases, we quote the  $1\text{-}\sigma$  error bars if the minimum of the  $1\text{-}\sigma$  error region is positive. Otherwise, we give the 90% confidence upper limit. While the *NuSTAR* observation taken on February 23 provides the only highly significant detection, the *Swift* observation with the highest count rate is the one that occurred during this *NuSTAR* observation. The only other possible evidence for activity from NuSTAR J163433–4738.7 occurred on February 22, when *NuSTAR* obtained a  $2.7\text{-}\sigma$  detection in FPMA; however, this is not confirmed by the FPMB data.

For *Chandra* ObsIDs 12529 and 12532 (from 2011) and ObsID 15625 (from 2013), we analyzed the ACIS data to search for a detection of NuSTAR J163433–4738.7. Based on an inspection, no sources are apparent in the 0.3–10 keV or 3–10 keV images. The source is  $7'$  and  $5'$  from the *Chandra* aimpoint for the 2011 ObsIDs. At these off-axis angles, the 90% encircled energy fraction (EEF) radii (for 4.5 keV photons) are  $7.''4$  and  $4.''5$  for ObsIDs 12529 and 12532, respectively. ObsID 15625 was a dedicated pointing with the target on-axis, and the 90% EEF radius is  $2''$ . For ObsID 12529, the largest number of 3–10 keV counts within any  $7.''4$ -radius circle inside the *NuSTAR* error region is three, and, after accounting for background, we calculate a 90% confidence upper limit

of  $<1.9 \times 10^{-4} \text{ s}^{-1}$  on the count rate (see Table 1). For ObsID 12532, the largest number of 3–10 keV counts within any  $4.''5$ -radius circle inside the *NuSTAR* error region is two, and the count rate limit is  $<2.1 \times 10^{-4} \text{ s}^{-1}$ . For ObsID 15625, the largest number of 3–10 keV counts within a  $2''$ -radius circle is two, and the count rate limit is  $<5.0 \times 10^{-4} \text{ s}^{-1}$ . This is higher than for the 2011 ObsIDs because of the lower exposure time.

We made *NuSTAR* light curves for ObsID 40014007001 with several different time binnings between 0.1 s and 5500 s (the approximate satellite orbital period) in the 3–10 keV and 3–79 keV bandpasses. The lack of any apparent variability in the 0.1–10 s light curves rules out flares or bursts that might prove the presence of a neutron star. Figure 2 shows the 3–10 keV orbit-by-orbit light curve for FPMA and FPMB combined. At the 5500 s binning, a  $\chi^2$  test shows that the 3–10 keV and 3–79 keV light curves are consistent with the source being constant over  $\sim 40$  ks.

Next, we extracted *NuSTAR* FPMA and FPMB spectra for ObsID 40014007001 and an XRT spectrum for ObsID 00080508001, which are the two observations with significant detections of NuSTAR J163433–4738.7. These were fitted jointly by minimizing the Cash (or  $C$ ) statistic (Cash 1979) using the XSPEC software package. The statistical quality of the spectrum is low, and it is well fit by a power-law, a blackbody, or a thermal bremsstrahlung model. In all three cases, we included absorption using Wilms, Allen & McCray (2000) abundances and Verner et al. (1996) cross-sections. The parameters are given in Table 2, and the errors quoted are 90% confidence for one parameter of interest,  $\Delta C = 2.7$ . Although we used the  $C$ -statistic to determine the parameters, we also calculated the reduced- $\chi^2$  values, and they appear in Table 2. This quantity is slightly smaller for the power-law model ( $\chi^2_\nu = 1.23$  for 8 degrees of freedom) compared to the blackbody model ( $\chi^2_\nu = 1.44$  for 8 dof) and the bremsstrahlung model ( $\chi^2_\nu = 1.29$  for 8 dof). However, given the small number of dof, the difference in  $\chi^2_\nu$  is not significant, and the steeply falling power-law index ( $\Gamma = 4.1^{+1.5}_{-1.0}$ ) suggests that the emission probably has a thermal origin. The spectrum for the blackbody model is shown in Figure 3.

We used the absorbed blackbody model and the count rates for *NuSTAR*, *Swift*, and *Chandra* given in Table 1 to calculate measurements or upper limits on the absorbed 3–10 keV flux, and these flux histories are shown in Figure 4. We also calculated the flux upper limits for the *Chandra* observations from 2011. The upper limits on the absorbed 3–10 keV fluxes are  $<1.6 \times 10^{-14} \text{ erg cm}^{-2} \text{ s}^{-1}$  and  $<2.5 \times 10^{-14} \text{ erg cm}^{-2} \text{ s}^{-1}$  for ObsIDs 12529 and 12532, respectively. Although the count rate limits are considerably lower for these ObsIDs compared to *Chandra* ObsID 15625, the flux limits are similar because of the different effective areas for the ACIS-I and ACIS-S instruments. The lowest *Chandra* upper limit indicates that the flux from this source changed by at least a factor of  $39^{+6}_{-4}$ , suggesting that it is a transient rather than simply being a highly variable X-ray source.

The *NuSTAR* error circle includes dozens of near-IR candidate counterparts brighter than  $K_s \sim 17$  (Vega magnitude system). In the NEWFIRM images from 2011, the brightest source in the error circle (2MASS J16343288–4738393) has  $K_s = 12.33 \pm 0.05$ ,  $H = 13.07 \pm 0.06$ , and  $J = 14.61 \pm 0.05$ , and the 2MASS magnitudes are consistent, indicating a possible lack of variability on long time scales. The next two brightest sources are 2MASS J16343362–4738479 at  $K_s = 13.31 \pm 0.05$

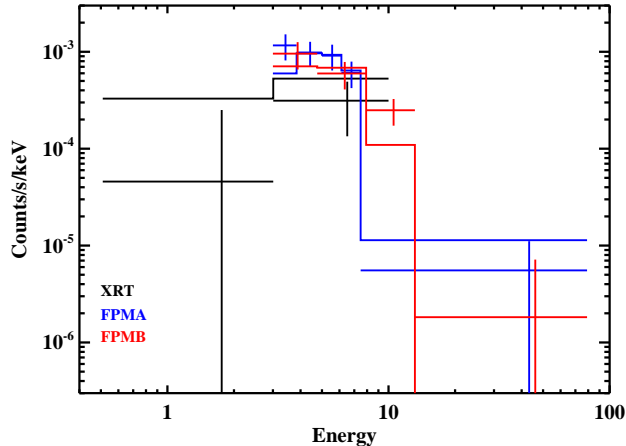


FIG. 3.— The *NuSTAR* and *Swift* energy spectra for NuSTAR J163433–4738.7 on 2013 Feb. 23–24 fitted with an absorbed blackbody model. The black, blue, and red data points and model lines are for XRT, FPMA, and FPMB, respectively. The errors on the data points are  $1\text{-}\sigma$  confidence.

and the *Spitzer*/GLIMPSE source G336.7870+00.0111 at  $K_s = 14.70 \pm 0.05$ . These may be more likely counterparts to NuSTAR J163433–4738.7 because they are relatively highly reddened ( $H - K_s = 2.35 \pm 0.08$  and  $1.34 \pm 0.08$ , respectively). However, we do not find any evidence that these sources are variable. For the two OSIRIS  $K_s$  images taken on April 3 and April 5, we performed aperture photometry for all the sources within the *NuSTAR* error circle, but did not find any variable sources.

The high time-resolution of the *NuSTAR* data allows for a search for a coherent signal with periods  $P \geq 4$  ms, covering the range expected for either an isolated rotation-powered pulsar, a binary, or a magnetar. Given the paucity of source counts in the observations listed in Table 1, we concentrate our attention on the ObsID 40014007001 data. Photon arrival times, adjusted for the *NuSTAR* clock drift, were corrected to the Solar System barycenter using JPL DE200 ephemeris and the *NuSTAR* derived source coordinates. We extracted photons in the 3–10 keV band from a  $30''$ -radius aperture centered on the source to optimize the signal-to-noise ratio. We searched for significant power from a coherent signal using an FFT sampled at the Nyquist frequency. The observation span was too short to consider the smearing of the pulse profile by spin-down of even the most highly energetic pulsar or the typical binary orbit period. The most significant signal found has a power of  $P = 35.18$ , corresponding to a probability of false detection of  $\wp = 0.8$  for  $2^{25}$  FFT search trials. We conclude that no pulsed X-ray signal is detected in from NuSTAR J163433–4738.7. After taking into account the local background, we place an upper limit on the pulse fraction at the  $3\text{-}\sigma$  confidence level of  $f_p < 36\%$  for a blind search for a sinusoidal signal  $P > 4$  ms.

#### 4. DISCUSSION

In considering the nature of NuSTAR J163433–4738.7, it is useful to estimate the X-ray luminosity that the source reached during its outburst. Although the distance is highly uncertain, the strong absorption either indicates a large distance through a region of the Galaxy with heavy extinction or absorption local to the source. The former is a strong possibility since the transient is in the direction

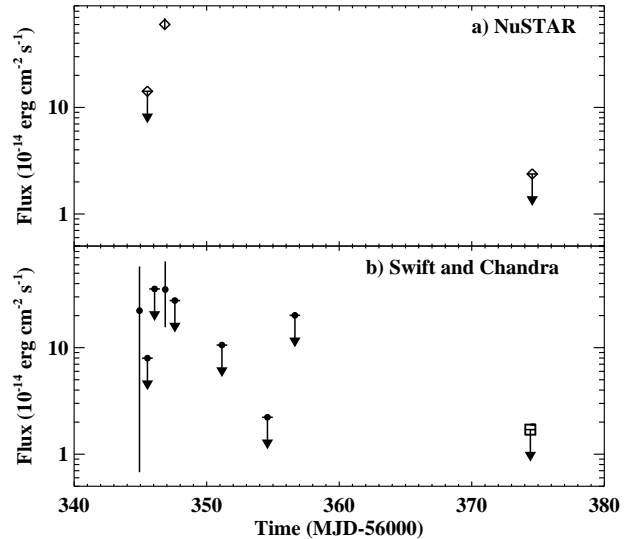


FIG. 4.— *NuSTAR*, *Swift*, and *Chandra* (marked with diamonds, filled circles, and a square, respectively) absorbed 3–10 keV fluxes for NuSTAR J163433–4738.7 assuming an absorbed blackbody with a column density of  $N_H = 9 \times 10^{22} \text{ cm}^{-2}$  and a temperature of 1.2 keV. The errors on the data points are  $1\text{-}\sigma$  significance, and the upper limits are 90% confidence.

of a region of the Galactic plane that is crowded with H II/molecular cloud regions. NuSTAR J163433–4738.7 is at  $l = 336.787^\circ$ ,  $b = +0.014^\circ$ , near a group of H II regions ( $336.732+0.072$ ,  $336.840+0.047$ ,  $336.9-0.1$ ,  $337.147-0.181$ , and  $337.3-0.1$ ) at a distance of  $10.9 \pm 0.2$  kpc (Georgelin et al. 1996; Russeil 2003)<sup>17</sup>. If the absorption is interstellar, then NuSTAR J163433–4738.7 is very likely beyond or in these molecular clouds, indicating a 2–10 keV unabsorbed luminosity limit of  $> 9 \times 10^{33} \text{ erg s}^{-1}$  for the blackbody spectrum and  $> 1.4 \times 10^{34} \text{ erg s}^{-1}$  for bremsstrahlung. In the following, we consider both possibilities: a distance of  $\sim 11$  kpc and a peak luminosity of  $\sim 10^{34} \text{ erg s}^{-1}$ ; and a smaller distance with absorption local to the source and a lower luminosity.

Considering the first possibility, a luminosity as high as  $10^{34} \text{ erg s}^{-1}$  would be unprecedented for at least two common types of X-ray sources in the Galaxy. Non-magnetic CVs have strong optical outbursts along with persistent and transient X-ray emission. However, extensive studies have shown that their X-ray luminosities are in the  $10^{29-32} \text{ erg s}^{-1}$  range (Baskill, Wheatley & Osborne 2005; Kuulkers et al. 2006). While magnetic CVs can reach higher luminosities (Kuulkers et al. 2006), they do not typically show outbursts. Secondly, active binaries, including RS CVn systems and low-mass flare stars, produce X-ray flares, some of which can last for about a day. Superflares that peak at X-ray luminosities of  $10^{32-33} \text{ erg s}^{-1}$  have been seen (Franciosini, Pallavicini & Tagliaferri 2001; Osten et al. 2007, 2010; Pandey & Singh 2012), and there have been flares that have released  $\gtrsim 10^{37} \text{ erg}$  (Franciosini, Pallavicini & Tagliaferri 2001). If NuSTAR J163433–4738.7 is at  $\sim 11$  kpc, then in addition to having a higher peak luminosity, the energy released is  $\gtrsim 4 \times 10^{38} \text{ erg}$ , which is based on the source being at its peak

<sup>17</sup> This kinematic distance is estimated using the mean Galactic rotation curve, and the uncertainty does not account for possible deviations relative to the mean.



luminosity for  $\gtrsim 40$  ks (see Figure 2). In summary, while we do not rule out the active binary possibility (and see below for further discussion on this topic), at the 11 kpc distance, the event seen by *NuSTAR* would need to be extreme for this explanation to be correct.

A source type that might be a good match to the NuSTAR J163433–4738.7 properties is the class of highly magnetic isolated neutron stars: magnetars. While these sources are best known for their very bright and brief ( $\sim 0.1$  s) X-ray and gamma-ray flares, they also have persistent but variable emission that can easily reach  $10^{34}$  erg s $^{-1}$  or higher (Woods & Thompson 2006). The X-ray spectrum of their persistent emission is dominated by a  $\sim 0.5$ – $1$  keV blackbody (Woods & Thompson 2006; Mori et al. 2013). Furthermore, as mentioned in § 1, magnetars are associated with regions where high-mass star formation is occurring, such as the Norma region, and there is a known magnetar, SGR 1627–41, that is  $\sim 0.2^\circ$  from NuSTAR J163433–4738.7. However, one possible counter-argument to the magnetar hypothesis is that magnetar periods of activity usually last for months. From Figure 4, the peak activity from NuSTAR J163433–4738.7 could not have lasted for more than  $\sim 1.5$  days based on the *Swift* upper limits, and *Chandra* and *NuSTAR* place a tight upper limit on activity  $\sim 3$  weeks after the *NuSTAR* detection. While we do not detect pulsations or  $\sim 0.1$  s flares, which would prove that the source is a magnetar, the pulsation search is limited by the statistical quality of the data, and flaring episodes are relatively rare.

A blackbody spectrum could also be produced from an optically thick accretion disk in a black hole binary. Most often, such sources show temperatures of  $\sim 1$  keV in the inner parts of their accretion disks when they are at high luminosity  $\gtrsim 10^{36-37}$  erg s $^{-1}$ , which would require a very large distance of  $\gtrsim 100$ – $300$  kpc for NuSTAR J163433–4738.7. Two sources (GRS 1758–258 and 1E 1740.7–2942) have shown fainter soft states, but their blackbody spectra have been at temperatures of 0.4 keV (Smith et al. 2001) and 0.7 keV (del Santo et al. 2005), which are lower than the 1.2 keV we see for NuSTAR J163433–4738.7. Also, transient black hole binaries typically have outbursts that last for several months, which is very different from NuSTAR J163433–4738.7.

Considering the second possibility that the source is closer than  $\sim 11$  kpc, then the high column density must be due to material local to the source. In fact, *INTEGRAL* has found a large number of obscured HMXBs that have column densities of  $10^{23-24}$  cm $^{-2}$  due to the wind from their supergiant companions (e.g. Walter et al. 2006). While many of the sources in this class are persistent, the Supergiant Fast X-ray Transients (SFXTs), which have some members that are also obscured HMXBs, have outbursts that can be as short as a few hours, and it would not be unusual for an SFXT to produce X-ray emission that lasts for half a day (Smith et al. 1998; Negueruela et al. 2006; Romano et al. 2011). However, a possible problem for this interpretation is that the spectrum of NuSTAR J163433–4738.7 is softer than usually seen for active SFXTs. During flares, IGR J17544–2619 and IGR J16479–4514 show a power-law spectrum with a photon index near  $\Gamma = 1.4$  (Rampy, Smith & Negueruela 2009; Sidoli et al. 2013). Somewhat softer spectra with  $\Gamma \sim 2.3$ – $2.4$  can be seen during periods of weaker activity, but their spectra are only as soft as NuSTAR J163433–4738.7 when they are in quiescence.

For CVs and active binaries, the challenge is to explain

a column density as high as we measure with absorption local to the systems. One possibility is that material expelled from the system (either related to the evolutionary state of the stellar component or related to the X-ray flaring event) could cause increased column density. However, for one active binary (AX Ari), the column density was seen to increase to  $N_{\text{H}} = 1.1 \times 10^{20}$  cm $^{-2}$  during an event (Franciosini, Pallavicini & Tagliaferri 2001), which is 2–3 orders of magnitude less than is required for NuSTAR J163433–4738.7. Variation in optical brightness of some stars has also been attributed to ejection of material from the stars. An example of this is a change of 0.5 magnitudes in the optical seen for KU Cyg (Tang et al. 2011). However, this would translate into a column density near  $10^{21}$  cm $^{-2}$ , which is still far less than is required.

The near-IR information that we have provides only weak constraints on the source type. Given that we were not able to identify the counterpart, we can only say that the source must be fainter (before and after the X-ray flare) than the brightest source in the error circle (2MASS J16343288–4738393). This corresponds to  $K_s > 12.3$ ,  $H > 13.1$ , and  $J > 14.6$ , and these values are consistent with all the possibilities for the source type discussed above. However, in the scenario where NuSTAR J163433–4738.7 is a nearby source with local absorption, these limits do provide some constraint. For example, a very nearby and luminous obscured HMXB is ruled out.

In summary, NuSTAR J163433–4738.7 is a fast X-ray transient, which has a thermal spectrum with relatively high temperature ( $kT = 1.2$  keV for a blackbody or 3.0 keV for bremsstrahlung). If its high column density is due to interstellar material, the source is probably distant ( $\gtrsim 11$  kpc), making the peak luminosity  $\gtrsim 10^{34}$  erg s $^{-1}$ . We discuss the origin of the flare and suggest that the most likely possibilities if the source is distant are an unusually bright flare from an active binary or a short outburst from a magnetar. We also consider the possibility that the source is closer and that the absorption is local to the source. More *NuSTAR* observations in the Galactic Plane will determine whether such transients are common and, hopefully, shed light on the nature of NuSTAR J163433–4738.7.

This work was supported under NASA Contract No. NNG08FD60C, and made use of data from the *NuSTAR* mission, a project led by the California Institute of Technology, managed by the Jet Propulsion Laboratory, and funded by the National Aeronautics and Space Administration. We thank the *NuSTAR* Operations, Software and Calibration teams for support with the execution and analysis of these observations. This research has made use of the *NuSTAR* Data Analysis Software (NuSTARDAS) jointly developed by the ASI Science Data Center (ASDC, Italy) and the California Institute of Technology (USA). RJA was supported by Gemini-CONICYT grant 32120009. FEB was supported by Basal-CATA PFB-06/2007 and CONICYT-Chile (through FONDECYT 1101024, gemini-conicyt 32120003, and Anillo ACT1101). LN wishes to acknowledge the Italian Space Agency (ASI) for financial support by ASI/INAF grant I/037/12/0-011/13. We thank Harvey Tananbaum for providing *Chandra* Director’s Discretionary Time for this project. This research has made use of the SIMBAD database, operated at CDS, Strasbourg, France.

TABLE 1  
X-RAY OBSERVATIONS AND COUNT RATES FOR NUSTAR J163433–4738.7

| Mission              | Instrument | ObsID       | Start Time (UT) | End Time (UT)   | Exposure (ks) | Count Rate <sup>a</sup>                |
|----------------------|------------|-------------|-----------------|-----------------|---------------|--|
| Observations in 2011 |            |             |                 |                 |               |  |
| <i>Chandra</i>       | ACIS-I     | 12529       | Jun 16, 6.97 h  | Jun 16, 12.53 h | 19.0          | <0.19                                  |
| <i>Chandra</i>       | ACIS-I     | 12532       | Jun 16, 23.85 h | Jun 17, 5.55 h  | 19.5          | <0.21                                  |
| Observations in 2013 |            |             |                 |                 |               |  |
| <i>Swift</i>         | XRT        | 00080509001 | Feb 21, 21.70 h | Feb 21, 23.13 h | 1.79          | 1.10 <sup>+1.77</sup> <sub>-1.06</sub> |
| <i>NuSTAR</i>        | FPMA       | 40014004001 | Feb 22, 7.77 h  | Feb 22, 17.52 h | 19.4          | 1.6 ± 0.6                              |
| "                    | FPMB       | "           | "               | "               | "             | <1.3                                   |
| <i>Swift</i>         | XRT        | 00080505001 | Feb 22, 11.88 h | Feb 22, 13.71 h | 1.92          | <0.54                                  |
| <i>Swift</i>         | XRT        | 00080506001 | Feb 23, 0.83 h  | Feb 23, 2.62 h  | 1.98          | <1.9                                   |
| <i>NuSTAR</i>        | FPMA       | 40014007001 | Feb 23, 14.52 h | Feb 24, 1.77 h  | 22.6          | 4.3 ± 0.6                              |
| "                    | FPMB       | "           | "               | "               | "             | 4.4 ± 0.9                              |
| <i>Swift</i>         | XRT        | 00080508001 | Feb 23, 20.20 h | Feb 23, 21.70 h | 1.97          | 2.2 <sup>+1.8</sup> <sub>-1.2</sub>    |
| <i>Swift</i>         | XRT        | 00080511001 | Feb 24, 13.72 h | Feb 24, 15.18 h | 1.75          | <0.82                                  |
| <i>Swift</i>         | XRT        | 00032728001 | Feb 28, 1.06 h  | Feb 28, 6.24 h  | 4.95          | <0.45                                  |
| <i>Swift</i>         | XRT        | 00032728002 | Mar 3, 9.00 h   | Mar 3, 19.06 h  | 4.85          | <0.12                                  |
| <i>Swift</i>         | XRT        | 00032728003 | Mar 5, 10.66 h  | Mar 5, 20.75 h  | 4.47          | <1.1                                   |
| <i>Chandra</i>       | ACIS-S     | 15625       | Mar 23, 8.30 h  | Mar 23, 11.91 h | 9.84          | <0.50                                  |
| <i>NuSTAR</i>        | FPMA       | 30001012002 | Mar 23, 8.52 h  | Mar 23, 18.35 h | 16.6          | <0.50                                  |
| "                    | FPMB       | "           | "               | "               | "             | <2.0                                   |

<sup>a</sup>Count rates in the 3–10 keV band in counts per ks. The errors given are 1- $\sigma$  and the upper limits are 90% confidence.

TABLE 2  
PARAMETERS FOR FITS TO *NuSTAR* AND *Swift*/XRT SPECTRA

| Model          | $N_{\text{H}}^{\text{a}}$                     | $\Gamma$ or $kT$                    | Flux <sup>b</sup>                   | C-statistic <sup>c</sup> | $\chi^2_{\nu}/\text{dof}$ |
|----------------|---|-------------------------------------|-------------------------------------|--------------------------|---------------------------|
| Power-law      | 28 <sup>+23</sup> <sub>-14</sub> <sup>d</sup> | 4.1 <sup>+1.5</sup> <sub>-1.0</sub> | 4 <sup>+18</sup> <sub>-2</sub>      | 9.5                      | 1.23/8                    |
| Blackbody      | 9 <sup>+15</sup> <sub>-7</sub>                | 1.2 ± 0.3 keV                       | 1.0 <sup>+0.8</sup> <sub>-0.3</sub> | 11.2                     | 1.44/8                    |
| Bremsstrahlung | 17 <sup>+17</sup> <sub>-9</sub>               | 3.0 <sup>+2.1</sup> <sub>-1.2</sub> | 1.6 <sup>+2.0</sup> <sub>-0.6</sub> | 9.8                      | 1.29/8                    |

<sup>a</sup>The column density in units of  $10^{22} \text{ cm}^{-2}$ .

<sup>b</sup>The 2–10 keV unabsorbed flux in units of  $10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$

<sup>c</sup>We fitted the spectra by minimizing the Cash statistic.

<sup>d</sup>All errors in this table are quoted at the 90% confidence level.

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