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D-brane Construction of the 5D NHEK Dual

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Abstract

Extremal but non-supersymmetric charged black holes with $SU(2)_L$ spin in IIB string theory compactified to five dimensions on $K^3 \times S^1$ are considered. These have a near-horizon or NHEK region with an enhanced $SL(2, R)_L$ conformal symmetry. It is shown that the NHEK geometry has a second, inequivalent, asymptotically flat extension in which the radius of the S^1 becomes infinite but the radius of the angular circles of $SU(2)_L$ orbits approach a constant. The asymptotic charges associated to the second solution identify it as a 5D D1-D5-Taub-NUT black string with certain nonzero worldvolume charge densities, temperatures and chemical potentials. The dual of the NHEK geometry is then identified as an IR limit of this wrapped brane configuration.

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1 Introduction

It has been conjectured [1] that extreme astrophysical Kerr black holes with spin J are dual to 2D conformal field theories with central charge $c = 12J$. In the real world, we cannot expect to know the exact form of the CFT, as that would be equivalent to knowing all the laws of physics to the Planck scale and beyond. Rather, a compelling match has been found between universal properties of the CFTs and universal properties of Kerr black holes. For a recent review see [2].

On the other hand it is nevertheless useful, in order to better understand the nature of the proposed real-world Kerr/CFT correspondence, to have a toy model in which the exact form of the CFT might be found. Such a model was recently constructed [3] by embedding certain 5D charged spinning black holes in string theory. At the maximal allowed value of the charge, the near horizon NHEK geometry is $\text{AdS}_3 \times S^2$, and standard string theoretic methods were used to identify the dual in terms of (a long string of) the wrapped D1-D5-Taub-NUT CFT. This result, as well as the properties of linear perturbations around maximality, all agree with expectations from the Kerr/CFT conjecture.

The issue of finite deviations from maximality, where the near-horizon geometry does not contain an AdS_3 factor, was not substantially addressed in [3]. In the present paper, we go beyond perturbative excursions from maximality and identify the dual for finite deformations as the IR limit of a wrapped D1-D5-Taub-NUT string with certain non-zero charge densities and temperatures. This is accomplished by constructing a 5D black string solution whose near-horizon geometry is locally identical to the 5D NHEK. The full black string geometry differs from the full black hole geometry in part because the angular circle does not grow in size at large radius. The microscopic D-brane configuration corresponding to the black string geometry is straightforward to read off from the asymptotic values of the electric and magnetic fields and is presented in table 1 below. The dual of the NHEK geometry, for generic charges and spins is thereby identified as an IR limit of this D-brane configuration.¹

¹ The precise nature of the flow to the IR (i.e. how charge densities are scaled etc.) can be determined from the radial dependence of the geometry. The nature of the IR limit here is clearly somewhat exotic because of the warping of the near-horizon AdS_3 , potentially related to a so-called dipole deformation of the IR CFT [4]. We do not address any of these interesting issues herein but hope to return to them elsewhere.

This paper is organized as follows. In section 2 we review the 5D charged black hole solutions described in [3], and introduce a slight generalization in which one of the charges and the asymptotic scalar are varied. In section 3 we construct new, asymptotically flat black string solutions of the same theory. The most general rotationally-invariant black string solution of 5D minimal supergravity does not have a warped near-horizon region [5]: a second $U(1)$ charge is needed for this purpose. Using the asymptotic values of the various fields at spatial infinity we compute and exhibit (see table 1) all the various magnetic charges, electric charge densities and ADM energy-momentum of the black string. This information enables us to identify the corresponding wrapped D-brane configuration. In section 4 we show that the near-horizon region of the 5D black holes of section 2 are a finite-temperature identification of the near-horizon region of the black strings of section 3, thereby identifying the specific collection of wrapped D-branes whose IR limit is dual to the 5D charged NHEK geometry.

2 5D extremal spinning black holes

In this section we describe non-supersymmetric, extremal, charged spinning black holes which arise in the compactification of IIB string theory to 5 dimensions on $K3 \times S^1$ and their NHEK limits.

2.1 The full solution

We consider the 6-dimensional Einstein-frame effective action

$$\mathcal{S}_{6B} = \frac{1}{8\pi^3} \int d^6x \sqrt{-g} \left(R - \frac{1}{12} (F_{(3)}^{RR})^2 \right). \quad (2.1)$$

This truncation of the low energy action without a dilaton is consistent when we make the additional self-duality restriction

$$F_{(3)}^{RR} = *F_{(3)}^{RR}. \quad (2.2)$$

We further specialize to black hole solutions of this action in a Kaluza-Klein (KK) compactification to $D = 5$ which carry only $SU(2)_L$ angular momentum J_L . These restrictions are made in order to illustrate the basic concepts in the simplest possible setting. More general solutions can be obtained by U-duality.

The KK compactification of (2.1) to 5 dimensions has one $U(1)$ gauge field descended from $F_{(3)}^{RR}$, a second KK $U(1)$ gauge field and a scalar parametrizing the radius of the S^1 . A careful presentation of the relation between 5D and 6D variables in this compactification can be found in [6, 7], but we mostly employ the simpler 6D variables throughout. The $J_R = 0$ extreme black hole solutions are labeled by three internal parameters: the mass (or J_L) and two charges. The asymptotic constant value of the scalar is a fourth external parameter. More general 5D black holes with more nonzero charges can be found in [8], whose six dimensional embedding can be found in [9].

We first review the subset of solutions studied in [3] characterized by the two parameters a and δ , and then present the simple two-parameter generalization. a and δ are related to

the mass M , graviphoton charge Q and spin J_L by

$$M = 6a^2 \cosh 2\delta \quad (2.3)$$

$$J_L = 4a^3 (c^3 + s^3), \quad (2.4)$$

$$Q = 4a^2 sc \quad (2.5)$$

where $c = \cosh \delta$ and $s = \sinh \delta$. These relations imply

$$Q^3 \leq J_L^2, \quad \left(\frac{M}{3} - \frac{Q}{2} \right)^2 \left(\frac{2M}{3} + 2Q \right) = J_L^2. \quad (2.6)$$

The metric is

$$\begin{aligned} ds_6^2 = & -\frac{(\hat{r}^2 + a^2(1 - 4c^2))}{\Sigma} d\hat{t}^2 + d\hat{u}^2 + \frac{4a^2 \sinh 2\delta}{\Sigma} d\hat{t}d\hat{u} + \Sigma \left(\frac{\hat{r}^2 d\hat{r}^2}{(\hat{r}^2 - a^2)^2} + \frac{d\theta^2}{4} \right) \\ & + \frac{\Sigma}{4} (d\hat{\psi}^2 + d\hat{\phi}^2 + 2 \cos \theta d\hat{\phi} d\hat{\psi}) + \frac{a^4}{\Sigma} (d\hat{\psi} + \cos \theta d\hat{\phi})^2 \\ & - \frac{4a^3}{\Sigma} ((c^3 + s^3) d\hat{t} + (s^2 c + c^2 s) d\hat{u}) (d\hat{\psi} + \cos \theta d\hat{\phi}) \end{aligned} \quad (2.7)$$

where $\Sigma \equiv \hat{r}^2 + a^2(1 + 4s^2)$ and

$$\hat{u} \sim \hat{u} + 2\pi m \quad (2.8)$$

is the coordinate of the unit-radius KK S^1 . Defining the one-form

$$A = \frac{2a^2 \sinh 2\delta}{\Sigma} \left(d\hat{t} - \frac{1}{2} a e^\delta (d\hat{\psi} + \cos \theta d\hat{\phi}) \right) \quad (2.9)$$

the RR three-form is given by the manifestly self-dual expression

$$F_{(3)}^{RR} = -dA \wedge (d\hat{u} + A) - *(dA \wedge (d\hat{u} + A)). \quad (2.10)$$

This two-parameter family of solutions is easily embedded in the larger four-parameter family, in which the two $U(1)$ charges and scalar field are separately varied. We simply deform the identification (2.25) by the parameters λ and R to

$$\hat{u} \sim \hat{u} + 2\pi m R \cosh \lambda, \quad \hat{t} \sim \hat{t} - 2\pi m R \sinh \lambda, \quad (2.11)$$

which does not introduce any singularities. To understand the asymptotic geometry we boost to the primed coordinates

$$\hat{u}' = \cosh \lambda \hat{u} + \sinh \lambda \hat{t}, \quad \hat{t}' = \cosh \lambda \hat{t} + \sinh \lambda \hat{u}, \quad (2.12)$$

which have the canonical identification

$$\hat{u}' \sim \hat{u}' + 2\pi m R, \quad \hat{t}' \sim \hat{t}'. \quad (2.13)$$

Hence the asymptotic geometry differs only by the radius R of the KK S^1 , previously set to unity.

2.2 The NHEK limit

In this subsection we describe the near-horizon limit, following [10, 11, 3]. Define

$$t = \frac{\Omega_L}{2a^2} \hat{t} \epsilon, \quad r = \frac{\hat{r}^2 - a^2}{\epsilon}, \quad y = 2\pi T_Q (\hat{u} + \Phi \hat{t}), \quad (2.14)$$

$$\psi = \hat{\psi} - \Omega_L \hat{t} - \frac{2J_L}{Q^2} (\hat{u} + \Phi \hat{t}), \quad \phi = \hat{\phi} \quad (2.15)$$

where

$$T_Q \equiv \frac{c^3 - s^3}{4\pi a s^2 c^2} = \frac{\sqrt{J_L^2 - Q^3}}{\pi Q^2}, \quad (2.16)$$

$$\Omega_L \equiv \frac{1}{a(c^3 - s^3)}, \quad (2.17)$$

$$\Phi \equiv \frac{c^2 s - s^2 c}{c^3 - s^3}. \quad (2.18)$$

The near-horizon metric is then the $\epsilon \rightarrow 0$ limit of (2.7)

$$\begin{aligned} \frac{12}{M} ds^2 &= -r^2 dt^2 + \frac{dr^2}{r^2} + \gamma(dy + rdt)^2 + \gamma(d\psi + \cos\theta d\phi)^2 \\ &+ 2\alpha\gamma(dy + rdt)(d\psi + \cos\theta d\phi) + d\theta^2 + \sin^2\theta d\phi^2 \end{aligned} \quad (2.19)$$

where the deformation parameters

$$\alpha = \frac{2 \cosh 2\delta}{1 + \cosh^2 2\delta}, \quad \gamma = 1 + \frac{1}{\cosh^2 2\delta}. \quad (2.20)$$

are related by $M\alpha\gamma = 12a^2$.

In terms of the $SL(2, R)_L \times SU(2)_R$ invariant forms,

$$\begin{aligned} \sigma_1 &= \cos\psi d\theta + \sin\theta \sin\psi d\phi \\ \sigma_2 &= -\sin\psi d\theta + \sin\theta \cos\psi d\phi \\ \sigma_3 &= d\psi + \cos\theta d\phi \\ w_{\pm} &= -e^{\mp y} r dt \mp e^{\mp y} dr/r \\ w_3 &= dy + r dt, \end{aligned} \quad (2.21)$$

the metric can be written

$$\frac{12}{M} ds^2 = -w_+ w_- + \gamma w_3^2 + \sigma_1^2 + \sigma_2^2 + \gamma \sigma_3^2 + 2\alpha\gamma w_3 \sigma_3. \quad (2.22)$$

The gauge field A reduces to

$$d\hat{u} + A = -\frac{a}{2} \tanh 2\delta (e^{\delta} \sigma_3 + e^{-\delta} w_3). \quad (2.23)$$

$F_{(3)}^{RR}$ follows from A via (2.10) as

$$F_{(3)}^{RR} = \frac{Q}{4} \left(\sigma_1 \wedge \sigma_2 \wedge \sigma_3 + \frac{1}{2} w_+ \wedge w_- \wedge w_3 + \operatorname{sech} 2\delta (\sigma_1 \wedge \sigma_2 \wedge w_3 + \frac{1}{2} w_+ \wedge w_- \wedge \sigma_3) \right) \quad (2.24)$$

The (y, ψ) identifications (2.11) are

$$y \sim y + 4\pi^2 T_R m, \quad \psi \sim \psi + 2\pi \Theta m + 4\pi n, \quad (2.25)$$

for any integers (m, n) . Here we have defined

$$T_R = RT_Q (\cosh \lambda - \Phi \sinh \lambda), \quad (2.26)$$

$$\Theta = -\frac{2J_L R}{Q^2} \cosh \lambda + R(\Omega_L + \frac{2\Phi J_L}{Q^2}) \sinh \lambda. \quad (2.27)$$

3 5D charged black strings

In this section we find 5D black string solutions of the same theory (2.1) which have the same near-horizon geometries as the 5D spinning black holes of the preceding section. The solutions are asymptotically flat in three, rather than four spatial directions and translationally invariant along the fourth.

3.1 Full solution and charges

A 5D black string is characterized by two magnetic charges obtained by integrating the two $U(1)$ field strengths over the S^2 surrounding the string. We will fix the KK magnetic charge to unity. This corresponds to having one Taub-NUT magnetic string.² The second RR magnetic charge, which we denote \mathcal{Q} below, counts both the number of $K3$ -wrapped D5-branes and parallel D1 branes which are equal due to the self-duality constraint we have imposed. The magnetic string worldvolume contains two conserved $U(1)$ currents arising from the two 5D bulk $U(1)$ gauge symmetries. The general black string solution has nonzero charge densities q_{KK} and q_{RR} along the string detected by long range electric fields at infinity. Hence we expect that the lowest-energy translationally and rotationally invariant configuration is described by the three internal parameters \mathcal{Q}, q_{KK} and q_{RR} . We also consider a momentum density along the string, but this is not an independent fourth parameter as it can always be eliminated by a boost.

Defining

$$\begin{aligned} \Omega_2 &= \sin \theta d\theta \wedge d\phi, & A_e &= -\frac{d\hat{t}}{\hat{r} + P} \\ H &= 1 + \frac{P}{\hat{r}} \end{aligned}$$

²Solutions with KK magnetic charge p are related to Z_p quotients of 5D black holes.

the action (2.1) has the two-parameter family of solutions

$$\begin{aligned}
ds^2 &= -H^{-2}d\hat{t}^2 + H^2(d\hat{r}^2 + \hat{r}^2(d\theta^2 + \sin^2\theta d\phi^2)) \\
&\quad + P^2\gamma[(d\hat{\psi} + \cos\theta d\phi + \alpha A_e)^2 + (d\hat{y} + \sqrt{1-\alpha^2}A_e)^2] \\
F_{(3)}^{RR} &= P^2 \tanh 2\delta \left(dA_e \wedge (d\hat{y} + A_e) + \Omega_2 \wedge (d\hat{\psi} + \cos\theta d\phi) \right) \\
&\quad - \frac{P^2 \tanh 2\delta}{\cosh 2\delta} \left(-dA_e \wedge (d\hat{\psi} + \cos\theta d\phi) + \Omega_2 \wedge (d\hat{y} - A_e) \right).
\end{aligned} \tag{3.1}$$

where α , γ are related to δ by (2.20). Regularity at the poles requires

$$\hat{\psi} \sim \hat{\psi} + 4\pi. \tag{3.2}$$

It can be verified that the RR 3-form is self-dual. At large \hat{r} , $H \rightarrow 1$ and the metric is

$$ds^2 = -d\hat{t}^2 + d\hat{r}^2 + \hat{r}^2(d\theta^2 + \sin^2\theta d\phi^2) + P^2\gamma d\hat{y}^2 + P^2\gamma(d\hat{\psi} + \cos\theta d\phi)^2. \tag{3.3}$$

This is locally a product of an S^1 parameterized by $\hat{\psi}$ with 5-dimensional Minkowski space and can be viewed as a Kaluza-Klein (KK) compactification to five dimensions.³

Note that the S^1 parameterized by $\hat{\psi}$ is fibered over the S^2 parameterized by (θ, ϕ) in a manner indicating there is a string extending in the \hat{y} direction carrying one unit of KK magnetic charge. This string also carries D1 and D5 charges Q_1 and Q_5 given by

$$Q_1 = Q_5 = \frac{1}{4\pi^2} \int_{S_3} F_{(3)}^{RR} = 4P^2 \tanh 2\delta \equiv \mathcal{Q}. \tag{3.4}$$

In addition the string carries electric charge densities. To see this we dimensionally reduce to five dimensions along the $\hat{\psi}$ circle, with the ansatz

$$ds_6^2 = \eta^{-\frac{1}{3}} ds_5^2 + \eta(d\hat{\psi} + \mathcal{A})^2 \tag{3.5}$$

which gives the 5D effective action

$$\mathcal{S}_{5B} = \frac{1}{2\pi^2} \int d^5x \sqrt{-g} \left\{ R - \frac{1}{3} \left(\frac{\partial\eta}{\eta} \right)^2 - \frac{1}{6} \eta^{\frac{2}{3}} (F_{(3)}^{RR})^2 - \frac{1}{4} \eta^{\frac{4}{3}} (\mathcal{F})^2 \right\} \tag{3.6}$$

where $\mathcal{F} = d\mathcal{A}$ and the KK scalar for the black string (2.7) is given by

$$\eta = P^2\gamma. \tag{3.7}$$

The 5D two-form arising from S^1 reduction of the three-form 6D RR field is proportional to the 5D dual of $F_{(3)}^{RR}$ due to the self duality condition, and so does not explicitly appear in (3.6). More details of the 6D to 5D reduction can be found in [6] and [7]. The KK gauge field can be read off of the $g_{\hat{\psi}\phi}$ and $g_{\hat{\psi}\hat{t}}$ terms in (3.1) and is

$$\mathcal{A} = \cos\theta d\phi + \alpha A_e. \tag{3.8}$$

³We expect that there is a more general solution in which the asymptotic radius of the KK S^1 can be varied arbitrarily, and that this is a special case in which the scalar sits at the attractor value throughout the geometry. As we are ultimately interested in the near-horizon behavior we will not consider this more general solution.

The first term gives the magnetic string charge

$$q_{TN} = -\frac{1}{4\pi} \int_{S^2} \mathcal{F} = 1 \quad (3.9)$$

while the second gives a radial electric field at large \hat{r}

$$\mathcal{F}_e = -\frac{\alpha}{\hat{r}^2} d\hat{t} \wedge d\hat{r}. \quad (3.10)$$

The proper density of KK electric charge on the string is hence

$$q_{KK} = -\frac{1}{\pi^2} \int_{S^2 \times R_y} \eta^{\frac{4}{3}} * \mathcal{F} = -\frac{\alpha}{\pi} (P^2 \gamma)^{\frac{4}{3}}. \quad (3.11)$$

where R_y is any interval along the y direction with unit proper length, and $*$ is the five dimensional Hodge dual. There is also an RR electric field at infinity. The corresponding charge density is

$$q_{RR} = \frac{1}{4\pi^2} \int_{S^2 \times R_y} F_{(3)}^{RR} = \frac{Q \operatorname{sech} 2\delta}{4\pi (P^2 \gamma)^{\frac{2}{3}}}. \quad (3.12)$$

This electric field is sourced by D1 branes wrapping the S^1 parameterized by $\hat{\psi}$ and smeared along the \hat{y} direction of the string, as well as smeared D5-branes wrapping $S^1 \times K3$. The proper densities of these objects along the string are equal and both given by q_{RR} . This string also has finite proper momentum and energy densities \mathcal{P} and \mathcal{E} . These are determined from the $\frac{1}{\hat{r}}$ corrections to the metric via the standard ADM formulae [12, 13, 14]. Using asymptotically Cartesian coordinates, the stress-energy tensor for a p -brane in d dimensional spacetime is

$$T_{ab} = \frac{1}{16\pi G_d} \oint d\Omega_{d-p-2} r^{d-p-2} n^i [\eta_{ab} (\partial_i h_c^c + \partial_i h_j^j - \partial_j h_i^i) - \partial_i h_{ab}] \quad (3.13)$$

where n^i is a radial unit vector in the transverse space, and $h_{\mu\nu} = g_{\mu\nu} - \eta_{\mu\nu}$. The labels $a, b = 0, 1 \dots p$ run over the world volume directions, while i, j denotes the transverse directions. For the black string solution (2.7), one first need to write the five dimensional Einstein frame metric in the canonical form, which amounts to the rescaling

$$\hat{t} = \eta^{-\frac{1}{6}} \tilde{t}, \quad \hat{r} = \eta^{-\frac{1}{6}} \tilde{r}, \quad \hat{y} = \eta^{-\frac{2}{3}} \tilde{y} \quad (3.14)$$

Rewriting the metric (2.7) in terms of tilded coordinates, and working out the stress tensor (3.13), one find the energy and momentum densities

$$\mathcal{E} = T_{\tilde{t}\tilde{t}} = \frac{8}{\pi} P (P^2 \gamma)^{\frac{1}{6}}, \quad \mathcal{P} = T_{\tilde{t}\tilde{y}} = -\frac{2}{\pi} (P^2 \gamma)^{2/3} \sqrt{1 - \alpha^2}. \quad (3.15)$$

One could transform everything to a boosted frame with $\mathcal{P} = 0$ but the resulting formulae are long and unilluminating.

In summary the supergravity solution (3.1) corresponds to the energy-momentum-density (3.15) brane configuration of the following table :

	Wrapping	6	7	8	9	u	\mathbb{R}_y
D5	\mathcal{Q}	X	X	X	X		X
D1	\mathcal{Q}						X
TN	1						X
D5	q_{RR}	X	X	X	X	X	
D1	q_{RR}						X
KK	q_{KK}						X

where \mathbb{R}_y denotes the common string direction, 6-7-8-9 are the $K3$ directions, the last three rows are proper wrapping densities and the last row denotes KK momentum around the S^1 .

3.2 Near-horizon limit

The near horizon limit is taken by defining

$$\hat{r} = \epsilon r, \quad \hat{t} = \frac{P^2 t}{\epsilon}, \quad (3.16)$$

$$\hat{\psi} = \psi + \alpha \left(y + \frac{Pt}{\epsilon} \right) \quad (3.17)$$

$$\hat{y} = \sqrt{1 - \alpha^2} \left(y + \frac{Pt}{\epsilon} \right) \quad (3.18)$$

Taking the limit $\epsilon \rightarrow 0$ gives the near-horizon metric

$$\frac{1}{P^2} ds^2 = -w_+ w_- + \gamma w_3^2 + \sigma_1^2 + \sigma_2^2 + \gamma \sigma_3^2 + 2\alpha \gamma w_3 \sigma_3 \quad (3.19)$$

with the identification

$$\psi \sim \hat{\psi} + 4\pi \quad (3.20)$$

and three-form

$$F_{(3)}^{RR} = \frac{\mathcal{Q}}{4} \left(\sigma_1 \wedge \sigma_2 \wedge \sigma_3 + \frac{1}{2} w_+ \wedge w_- \wedge w_3 + \text{sech} 2\delta \left(\sigma_1 \wedge \sigma_2 \wedge w_3 + \frac{1}{2} w_+ \wedge w_- \wedge \sigma_3 \right) \right) \quad (3.21)$$

This general near-horizon geometry is fully described by two parameters which can be taken to be δ and P . The asymptotically flat solution on the other hand has a three-parameter generalization in which q_{KK} and q_{RR} are separately varied. These all reduce, however, in the near-horizon limit to the same two parameter family given above. This corresponds, in the dual picture, to radial RG flows which differ in the UV but reach the same IR fixed point. This observation agrees with the analysis of [5], where general black string solutions in the minimal 5D supergravity with a single $U(1)$ were studied. In that case, the radial electric field is always scaled away in the near-horizon region,⁴ which is therefore always AdS_3 .⁵ We see here that in order to get a rotationally-invariant near-horizon warped AdS_3 two $U(1)$ s are required.

⁴For rotating black strings, a different decoupling limit was discussed in [15], under which there is also a warped horizon, but the $SU(2)$ symmetry is broken.

⁵The fact that one of the two electric fields scales away in the near-horizon region is related to the fact that the near horizon $SL(2, R)_L \times U(1)_R$ isometry contains a left but not a right scaling symmetry.

4 NHEK = black string near-horizon

The near-horizon black string geometry (3.19), (3.21) is locally the same as the NHEK geometry (2.22), (2.24) with the identification

$$M = 12P^2. \tag{4.1}$$

Globally the NHEK geometry also has the identification (2.25) of the y coordinate. In terms of the hatted coordinates of the full black string solution (3.1) this identification is

$$\hat{y} \sim \hat{y} + 4\pi^2 T_R \sqrt{1 - \alpha^2} m, \quad \hat{\psi} \sim \hat{\psi} + 2\pi(2\alpha\pi T_R + \Theta)m + 4\pi n. \tag{4.2}$$

In other words, the 5D black string is wrapped around a compactified \hat{y} circle and has boundary conditions twisted by a rotation along $\hat{\psi}$.

In conclusion, the dual of the 5D extreme spinning black holes is identified as the low energy limit of the D-brane configuration in table 1 wrapped on a circle with $\hat{\psi}$ -twisted boundary conditions.

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References

- [1] M. Guica, T. Hartman, W. Song, A. Strominger, “The Kerr/CFT Correspondence,” *Phys. Rev.* **D80**, 124008 (2009). [arXiv:0809.4266 [hep-th]].
- [2] I. Bredberg, C. Keeler, V. Lysov, A. Strominger, “Cargese Lectures on the Kerr/CFT Correspondence,” [arXiv:1103.2355 [hep-th]].
- [3] M. Guica and A. Strominger, “Microscopic Realization of the Kerr/CFT Correspondence,” arXiv:1009.5039 [hep-th].
- [4] A. Bergman, O. J. Ganor, “Dipoles, twists and noncommutative gauge theory,” *JHEP* **0010**, 018 (2000). [hep-th/0008030].
- [5] G. Compere, S. de Buyl, S. Stotyn, A. Virmani, “A General Black String and its Microscopics,” *JHEP* **1011**, 133 (2010). [arXiv:1006.5464 [hep-th]].
- [6] E. Cremmer, B. Julia, H. Lu and C. N. Pope, “Dualisation of dualities. I,” *Nucl. Phys. B* **523**, 73 (1998) [arXiv:hep-th/9710119].
- [7] M. J. Duff, H. Lu and C. N. Pope, “AdS(3) x S**3 (un)twisted and squashed, and an O(2,2,Z) multiplet of dyonic strings,” *Nucl. Phys. B* **544**, 145 (1999) [arXiv:hep-th/9807173].

- [8] M. Cvetič and D. Youm, “General rotating five-dimensional black holes of toroidally compactified heterotic string,” Nucl. Phys. B **476**, 118 (1996) [arXiv:hep-th/9603100].
- [9] S. Giusto, S. D. Mathur and A. Saxena, “Dual geometries for a set of 3-charge microstates,” Nucl. Phys. B **701**, 357 (2004) [arXiv:hep-th/0405017].
- [10] O. J. C. Dias, R. Emparan and A. Maccarrone, “Microscopic theory of black hole superradiance,” Phys. Rev. D **77**, 064018 (2008) [arXiv:0712.0791 [hep-th]].
- [11] I. Bredberg, T. Hartman, W. Song and A. Strominger, “Black Hole Superradiance From Kerr/CFT,” JHEP **1004**, 019 (2010) [arXiv:0907.3477 [hep-th]].
- [12] J. X. Lu, “ADM masses for black strings and p-branes,” Phys. Lett. B **313**, 29 (1993) [arXiv:hep-th/9304159].
- [13] R. C. Myers, “Stress tensors and Casimir energies in the AdS/CFT correspondence,” Phys. Rev. D **60**, 046002 (1999) [arXiv:hep-th/9903203].
- [14] T. Harmark and N. A. Obers, “General definition of gravitational tension,” JHEP **0405**, 043 (2004) [arXiv:hep-th/0403103].
- [15] G. Compere, W. Song and A. Virmani, “Microscopics of Extremal Kerr from Spinning M5 Branes,” arXiv:1010.0685 [hep-th].