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## SHORT GRB 130603B: DISCOVERY OF A JET BREAK IN THE OPTICAL AND RADIO AFTERGLOWS, AND A MYSTERIOUS LATE-TIME X-RAY EXCESS

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### ABSTRACT

We present radio, optical/NIR, and X-ray observations of the afterglow of the short-duration GRB 130603B, and uncover a break in the radio and optical bands at  $\approx 0.5$  d after the burst, best explained as a jet break with an inferred jet opening angle of  $\approx 4-8^\circ$ . GRB 130603B is only the third short GRB with a radio afterglow detection to date, and the first time that a jet break is evident in the radio band. We model the temporal evolution of the spectral energy distribution to determine the burst explosion properties and find an isotropic-equivalent kinetic energy of  $\approx (0.6-1.7) \times 10^{51}$  erg and a circumburst density of  $\approx 5 \times 10^{-3}-30$  cm<sup>-3</sup>. From the inferred opening angle of GRB 130603B, we calculate beaming-corrected energies of  $E_\gamma \approx (0.5-2) \times 10^{49}$  erg and  $E_K \approx (0.1-1.6) \times 10^{49}$  erg. Along with previous measurements and lower limits we find a median opening angle of  $\approx 10^\circ$ . Using the all-sky observed rate of 10 Gpc<sup>-3</sup> yr<sup>-1</sup>, this implies a true short GRB rate of  $\approx 20$  yr<sup>-1</sup> within 200 Mpc, the Advanced LIGO/VIRGO sensitivity range for neutron star binary mergers. Finally, we uncover evidence for significant excess emission in the X-ray afterglow of GRB 130603B at  $\gtrsim 1$  d and conclude that the additional energy component could be due to fall-back accretion or spin-down energy from a magnetar formed following the merger.

### 1. INTRODUCTION

The broad-band afterglows of short-duration gamma-ray bursts (GRBs;  $T_{90} < 2$  sec; Kouveliotou et al. 1993) provide a unique opportunity to study the basic properties of these events: their energetics, circumburst densities, and jet opening angles. While the energy scales and densities provide fundamental insight to the explosion physics and progenitors, the geometry provides direct information on the population's true energy scale and event rate. The most likely progenitors, neutron star-neutron star (NS-NS) and/or neutron star-black hole (NS-BH) mergers (Eichler et al. 1989; Narayan et al. 1992; Fong et al. 2010; Berger 2010; Fong & Berger 2013), are the premier candidates of gravitational waves for Advanced LIGO/VIRGO. Thus, an inference on the opening angle distribution will aid our expectations for coincident electromagnetic-gravitational wave detections.

Over the past  $\sim 9$  years, the *Swift* satellite (Gehrels et al. 2004) has discovered 77 short GRBs, with a median  $\gamma$ -ray fluence of  $\approx 2 \times 10^{-7}$  erg cm<sup>-2</sup> (15–150 keV; Gehrels et al. 2008; Nysewander et al. 2009). Coupled with redshift measurements from their host galaxies, this has demonstrated a range of isotropic-equivalent  $\gamma$ -ray energies of  $E_{\gamma,\text{iso}} \approx 10^{48}-10^{52}$  erg (Berger 2007; Nysewander et al. 2009). Temporal

monitoring of their broad-band afterglows has led to a similarly broad range of inferred isotropic-equivalent kinetic energies,  $E_{K,\text{iso}} \approx 10^{48}-10^{52}$  erg, circumburst densities of  $\lesssim 1$  cm<sup>-3</sup> (Soderberg et al. 2006; Panaitescu 2006; Stratta et al. 2007; Perley et al. 2009; Berger 2010; Fong et al. 2011), and opening angles of  $\gtrsim 3-25^\circ$  (Fox et al. 2005; Grupe et al. 2006; Soderberg et al. 2006; Fong et al. 2012; Margutti et al. 2012a; Nicuesa Guelbenzu et al. 2012; Berger et al. 2013; Sakamoto et al. 2013).

Collimation in GRBs is determined from temporal steepenings in afterglow light curves, termed “jet breaks”, which are expected to be achromatic (Sari et al. 1999; Rhoads 1999; Panaitescu 2005). Jet breaks in the light curves of long GRBs translate to an opening angle distribution with a range of  $\theta_j \approx 2-25^\circ$  and a median of  $\approx 7^\circ$ , leading to beaming-corrected energies of  $E_\gamma = [1 - \cos(\theta_j)]E_{\gamma,\text{iso}} \approx 10^{51}$  erg (Bloom et al. 2003; Frail et al. 2001; Friedman & Bloom 2005; Kocevski & Butler 2008; Racusin et al. 2009). For short GRBs, the search for jet breaks has been more challenging, primarily due to the intrinsic faintness of their afterglows, a direct reflection of their low energy scales and circumburst densities. There are only three robust cases of jet breaks for short GRBs based on well-sampled light curves thus far (GRB 051221A: Soderberg et al. 2006, GRB 090426: Nicuesa Guelbenzu et al. 2011; GRB 111020A: Fong et al. 2012), with  $\theta_j \approx 3-8^\circ$ . The non-detection of jet breaks to  $\gtrsim 1$  day after the burst has also provided lower limits on the jet opening angles of  $\gtrsim 3-25^\circ$  (Fox et al. 2005; Grupe et al. 2006; Fong et al. 2012; Coward et al. 2012; Margutti et al. 2012a; Berger et al. 2013), suggesting that short GRBs may have wider jets than their long-duration counterparts.

In addition to an inference on the opening angles, afterglows can also provide unique constraints on the energy scales and circumburst densities through multi-wavelength detections and modeling of the spectral energy distributions. However, of the 77 *Swift* short GRBs detected to present, only two have been detected in the radio band (GRB 050724A:

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Berger et al. 2005; GRB 051221A: Soderberg et al. 2006). Both of these events have inferred densities of  $\approx 10^{-2} \text{ cm}^{-3}$  and isotropic-equivalent energies of  $\approx 10^{51} \text{ erg}$ . The upgrade of the VLA with a ten-fold increase in sensitivity (Perley et al. 2011) provides a promising route to increased radio afterglow detections, and thus substantially tighter constraints on these properties.

Temporal afterglow information, particularly in the X-ray band, has also revealed cases of anomalous behavior that do not fit with the standard synchrotron picture of afterglows. For short GRBs, there have been observed cases of flares at  $\lesssim 1000 \text{ s}$  after the burst (Grupe et al. 2006; Margutti et al. 2011), shallow decays attributed to energy injection (e.g., Soderberg et al. 2006), putative early plateaus attributed to the spin-down power of a hyper-massive and highly-magnetized neutron star remnant (Rowlinson et al. 2010, 2013) and two cases of late-time X-ray re-brightenings on  $\sim \text{day}$  timescales (GRB 050724A: Grupe et al. 2006; GRB 080503: Perley et al. 2009).

Recently, the short-duration GRB 130603B sparked much interest, because its bright optical afterglow enabled the first afterglow spectrum of a short GRB and thus an unambiguous redshift determination of  $z = 0.3565 \pm 0.0002$  (Cucchiara et al. 2013; de Ugarte Postigo et al. 2013). It also led to the first claimed detection of a “kilonova” associated with a short GRB (Berger et al. 2013; Tanvir et al. 2013), providing direct evidence for a compact object binary progenitor. Here, we present and model the broad-band afterglow of GRB 130603B. In Section 2, we present the X-ray, optical/NIR and radio data sets. In Section 3, we model the evolution of the spectral energy distribution (SED) and constrain the burst explosion properties. In Section 4, we discuss the steepening in the radio and optical light curves, best explained as a jet break. In Section 5, we investigate several possibilities to explain excess X-ray emission at  $\gtrsim 1 \text{ day}$ . Finally in Section 6, we discuss GRB 130603B in the context of the short GRB population, and investigate implications for the energy scale and event rate.

Unless otherwise noted, all magnitudes in this paper are in the AB system and are corrected for Galactic extinction in the direction of the burst using  $E(B-V) = 0.02 \text{ mag}$  (Schlegel et al. 1998; Schlafly & Finkbeiner 2011). Reported uncertainties correspond to 68% confidence. We employ a standard  $\Lambda$ CDM cosmology with  $\Omega_M = 0.27$ ,  $\Omega_\Lambda = 0.73$ , and  $H_0 = 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$ .

## 2. OBSERVATIONS

GRB 130603B was detected on 2013 June 3 at 15:49:14 UT by the Burst Alert Telescope (BAT; Barthelmy et al. 2005) on-board the *Swift* satellite (Melandri et al. 2013), and by *Konus-Wind* (Golenetskii et al. 2013). *Swift*/BAT localized the burst to RA=11<sup>h</sup>28<sup>m</sup>53.2<sup>s</sup> and Dec=+17°03′48.2″ (J2000) with 1.0′ accuracy (90% containment; Barthelmy et al. 2013). The  $\gamma$ -ray emission consists of a single pulse with a duration of  $T_{90} = 0.18 \pm 0.02 \text{ s}$  (15–150 keV; Barthelmy et al. 2013), a fluence of  $f_\gamma = (6.6 \pm 0.7) \times 10^{-6} \text{ erg cm}^{-2}$  (20–10<sup>4</sup> keV; Golenetskii et al. 2013), and a peak energy of  $E_{\text{pk}} = 660 \pm 100 \text{ keV}$  (Golenetskii et al. 2013). The spectral lags are  $0.6 \pm 0.7 \text{ ms}$  (15–25 keV to 50–100 keV) and  $-2.5 \pm 0.7 \text{ ms}$  (25–50 keV to 100–350 keV; Norris et al. 2013). The combination of the duration, high peak energy, and zero (or negative) lag unambiguously classifies GRB 130603B as a short-hard burst. At  $z = 0.3565$  (Cucchiara et al. 2013; de Ugarte Postigo et al. 2013), the isotropic-equivalent gamma-ray energy is  $E_{\gamma,\text{iso}} \approx$

$2.1 \times 10^{51} \text{ erg}$  (20–10<sup>4</sup> keV, rest-frame).

### 2.1. X-rays

Observations with the X-ray Telescope (XRT; Burrows et al. 2005) on-board *Swift* commenced at  $\delta t = 59 \text{ s}$  (where  $\delta t$  is the time after the BAT trigger) and revealed a fading, uncatalogued X-ray source with a UVOT-enhanced position of RA=11<sup>h</sup>28<sup>m</sup>48.15<sup>s</sup> and Dec=+17°04′16.9″ (J2000) and an uncertainty of 1.4″ radius (90% containment; Goad et al. 2007; Evans et al. 2007, 2009). The source faded below the XRT detection threshold by  $\delta t \approx 2 \text{ d}$ . We analyze the XRT data using the HEASOFT package (v.6.13) and relevant calibration files. To generate a count-rate light curve, we apply standard filtering and screening criteria (see Margutti et al. 2012b), ensuring a minimum signal-to-noise ratio of 4 for each temporal bin.

We initiated a target-of-opportunity program on *XMM-Newton* (ID: 072257, PI: Fong) with the European Photon Imaging Camera (EPIC-PN) and obtained observations at  $\delta t \approx 2.7 \text{ d}$  and  $\approx 6.5 \text{ d}$  with net exposure times of 18.5 ks and 30.2 ks, respectively. We analyze these data using standard routines in the Scientific Analysis System (SAS) and detect a fading source coincident with the XRT position. We find  $\approx 60$  counts ( $\sim 4\sigma$ ) in the first observation in a 30″-radius aperture, and  $\approx 10$  counts ( $3\sigma$ ) in the second observation in a 15″-radius aperture, where the aperture size is adjusted to maximize the signal-to-noise ratio.

To determine the flux calibration, we model the XRT data with an absorbed single power-law spectrum, using the Galactic neutral Hydrogen absorption column ( $N_{\text{H,MW}} = 1.93 \times 10^{20} \text{ cm}^{-2}$ ; Kalberla et al. 2005). We fit for the photon index,  $\Gamma$  and the intrinsic Hydrogen absorption column ( $N_{\text{H,int}}$  at  $z = 0.3565$ ). Using all of the available XRT data, we find  $N_{\text{H,int}} = (2.4 \pm 0.4) \times 10^{21} \text{ cm}^{-2}$  and  $\Gamma = 2.2 \pm 0.1$ . We apply these spectral parameters to the XRT and XMM data. The resulting fluxes from XMM are listed in Table 1.

### 2.2. Optical

Subsequent to the discovery of the X-ray afterglow (Melandri et al. 2013), ground-based optical/NIR observations began at  $\delta t \approx 2.7 \text{ hr}$  to search for an optical counterpart (Berger et al. 2013; Cucchiara et al. 2013; de Ugarte Postigo et al. 2013; Tanvir et al. 2013). We initiated four sets of *r*-band observations of GRB 130603B using instruments on the twin Magellan 6.5-m telescopes, spanning  $\delta t \approx 8.1 \text{ hr}$  to  $\approx 37 \text{ d}$  (Table 1). The description of these observations and the optical afterglow discovery are provided in Berger et al. (2013). In addition, we obtained a set of *r*-band observations with MMTCam on the 6.5-m Multi-Mirror Telescope (MMT) and processed these data using standard procedures in IRAF. All of the published optical/NIR afterglow photometry (Berger et al. 2013; Cucchiara et al. 2013; de Ugarte Postigo et al. 2013; Tanvir et al. 2013), are listed in Table 1.

### 2.3. Radio

We observed the position of GRB 130603B with the Karl G. Jansky Very Large Array (VLA; Program 13A-046, PI: Berger) starting at  $\delta t \approx 8.8 \text{ hr}$  at a mean frequency of 5.8 GHz (upper and lower side-bands centered at 6.7 GHz and 4.9 GHz) using 3C286 and J1120-1420 for bandpass/flux and gain calibration, respectively. We follow standard procedures in the Astronomical Image Processing System (AIPS);

**Table 1**  
GRB 130603B Afterglow Photometry

| $\delta t$<br>(d) | Exposure Time<br>(hr) | Telescope         | Instrument | Band     | $F_\nu$<br>( $\mu\text{Jy}$ ) | $\sigma_\nu$<br>( $\mu\text{Jy}$ ) | References |
|-------------------|-----------------------|-------------------|------------|----------|-------------------------------|------------------------------------|------------|
| <i>X-rays</i>     |                       |                   |            |          |                               |                                    |            |
| 2.69              | 5.14                  | <i>XMM-Newton</i> | EPIC-pn    | 1 keV    | $2.44 \times 10^{-3}$         | $4.59 \times 10^{-4}$              | This work  |
| 6.50              | 8.38                  | <i>XMM-Newton</i> | EPIC-pn    | 1 keV    | $8.25 \times 10^{-4}$         | $5.03 \times 10^{-4}$              | This work  |
| <i>Optical</i>    |                       |                   |            |          |                               |                                    |            |
| 0.008             | 0.05                  | Swift             | UVOT       | <i>V</i> | < 199.5                       |                                    | 1          |
| 0.089             | 1.42                  | Swift             | UVOT       | <i>V</i> | < 53.95                       |                                    | 1          |
| 0.24              | 0.50                  | NOT               | MOS        | <i>r</i> | 12.59                         | 0.23                               | 2          |
| 0.25              | 0.25                  | WHT               | ACAM       | <i>z</i> | 25.35                         | 1.36                               | 1          |
| 0.27              | 0.25                  | WHT               | ACAM       | <i>i</i> | 16.44                         | 0.88                               | 1          |
| 0.28              | 0.50                  | CAHA              | DLR-MKIII  | <i>V</i> | 8.32                          | 0.73                               | 1          |
| 0.29              | 0.008                 | GTC               | OSIRIS     | <i>r</i> | 10.96                         | 0.20                               | 2          |
| 0.29              | 0.25                  | WHT               | ACAM       | <i>g</i> | 6.31                          | 0.34                               | 2          |
| 0.33              | 0.40                  | Gemini-South      | GMOS       | <i>g</i> | 5.30                          | 0.19                               | 3          |
| 0.34              | 0.17                  | Magellan/Baade    | IMACS      | <i>r</i> | 8.64                          | 0.14                               | 4          |
| 0.37              | 0.40                  | Gemini-South      | GMOS       | <i>i</i> | 12.25                         | 1.18                               | 3          |
| 0.60              | 0.14                  | Gemini-North      | GMOS       | <i>z</i> | 6.54                          | 0.18                               | 2          |
| 0.60              | 0.19                  | UKIRT             | WFCAM      | <i>K</i> | 13.68                         | 1.32                               | 2          |
| 0.61              | 0.14                  | Gemini-North      | GMOS       | <i>i</i> | 4.53                          | 0.12                               | 2          |
| 0.61              | 0.19                  | UKIRT             | WFCAM      | <i>J</i> | 9.29                          | 1.12                               | 2          |
| 0.62              | 0.14                  | Gemini-North      | GMOS       | <i>r</i> | 2.88                          | 0.08                               | 2          |
| 0.62              | 0.14                  | Gemini-North      | GMOS       | <i>g</i> | 1.60                          | 0.06                               | 2          |
| 1.30              | 0.15                  | Gemini-South      | GMOS       | <i>r</i> | < 0.30                        |                                    | 3          |
| 1.30              | 0.15                  | Gemini-South      | GMOS       | <i>i</i> | < 0.58                        |                                    | 3          |
| 1.34              | 0.33                  | Magellan/Baade    | IMACS      | <i>r</i> | < 0.46                        |                                    | 4          |
| 1.59              | 0.17                  | Gemini-North      | GMOS       | <i>g</i> | < 0.19                        |                                    | 2          |
| 1.60              | 0.17                  | Gemini-North      | GMOS       | <i>r</i> | 0.21                          | 0.05                               | 2          |
| 1.61              | 0.17                  | Gemini-North      | GMOS       | <i>i</i> | < 0.48                        |                                    | 2          |
| 1.61              | 0.39                  | UKIRT             | WFCAM      | <i>J</i> | < 1.25                        |                                    | 2          |
| 1.62              | 0.17                  | Gemini-North      | GMOS       | <i>z</i> | < 1.00                        |                                    | 2          |
| 2.32              | 0.37                  | VLT               | HAWK-I     | <i>J</i> | < 1.25                        |                                    | 2          |
| 3.26              | 0.17                  | GTC               | OSIRIS     | <i>r</i> | < 0.33                        |                                    | 2          |
| 4.26              | 0.17                  | GTC               | OSIRIS     | <i>r</i> | < 0.23                        |                                    | 2          |
| 7.30              | 0.37                  | VLT               | HAWK-I     | <i>J</i> | < 1.45                        |                                    | 2          |
| 8.23              | 0.33                  | TNG               | DOLoRes    | <i>r</i> | < 1.15                        |                                    | 1          |
| 8.25              | 0.33                  | TNG               | DOLoRes    | <i>i</i> | < 0.52                        |                                    | 1          |
| 8.41              | 0.33                  | Magellan/Baade    | IMACS      | <i>r</i> | < 0.40                        |                                    | This work  |
| 9.37              | 0.62                  | HST               | ACS        | F606W    | < 0.03                        |                                    | 4          |
| <sup>a</sup> 9.41 | 0.73                  | HST               | WFC3       | F160W    | 0.17                          | 0.03                               | 2, 4       |
| 21.26             | 0.67                  | TNG               | DOLoRes    | <i>r</i> | < 1.49                        |                                    | 1          |
| 21.29             | 0.67                  | TNG               | DOLoRes    | <i>i</i> | < 0.93                        |                                    | 1          |
| 21.52             | 0.28                  | MMT               | MMTCam     | <i>r</i> | < 1.91                        |                                    | This work  |
| 29.57             | 0.73                  | HST               | WFC3       | F160W    | < 0.10                        |                                    | 4          |
| 37.34             | 0.75                  | Magellan/Clay     | LDSS3      | <i>r</i> | < 0.40                        |                                    | This work  |
| <i>Radio</i>      |                       |                   |            |          |                               |                                    |            |
| 0.37              | 2.00                  | VLA               |            | 4.9 GHz  | 125.1                         | 14.4                               | This work  |
| 0.37              | 2.00                  | VLA               |            | 6.7 GHz  | 118.6                         | 9.1                                | This work  |
| 1.43              | 1.00                  | VLA               |            | 4.9 GHz  | < 56.7                        |                                    | This work  |
| 1.43              | 1.00                  | VLA               |            | 6.7 GHz  | 64.9                          | 15.2                               | This work  |
| 1.44              | 1.00                  | VLA               |            | 21.8 GHz | < 50.0                        |                                    | This work  |
| 4.32              | 2.00                  | VLA               |            | 4.9 GHz  | < 51.0                        |                                    | This work  |
| 4.32              | 2.00                  | VLA               |            | 6.7 GHz  | < 25.8                        |                                    | This work  |
| 84.31             | 1.00                  | VLA               |            | 4.9 GHz  | < 68.6                        |                                    | This work  |
| 84.31             | 1.00                  | VLA               |            | 6.7 GHz  | < 33.7                        |                                    | This work  |

**Note.** — All upper limits correspond to  $3\sigma$  confidence. Optical flux densities are corrected for Galactic extinction in the direction of the burst, but are not corrected for extinction in the rest-frame of the burst.

<sup>a</sup> The reported flux and uncertainty are of the claimed kilonova detection (Berger et al. 2013; Tanvir et al. 2013).

**References:** (1) de Ugarte Postigo et al. 2013; (2) Tanvir et al. 2013; (3) Cucchiara et al. 2013; (4) Berger et al. 2013

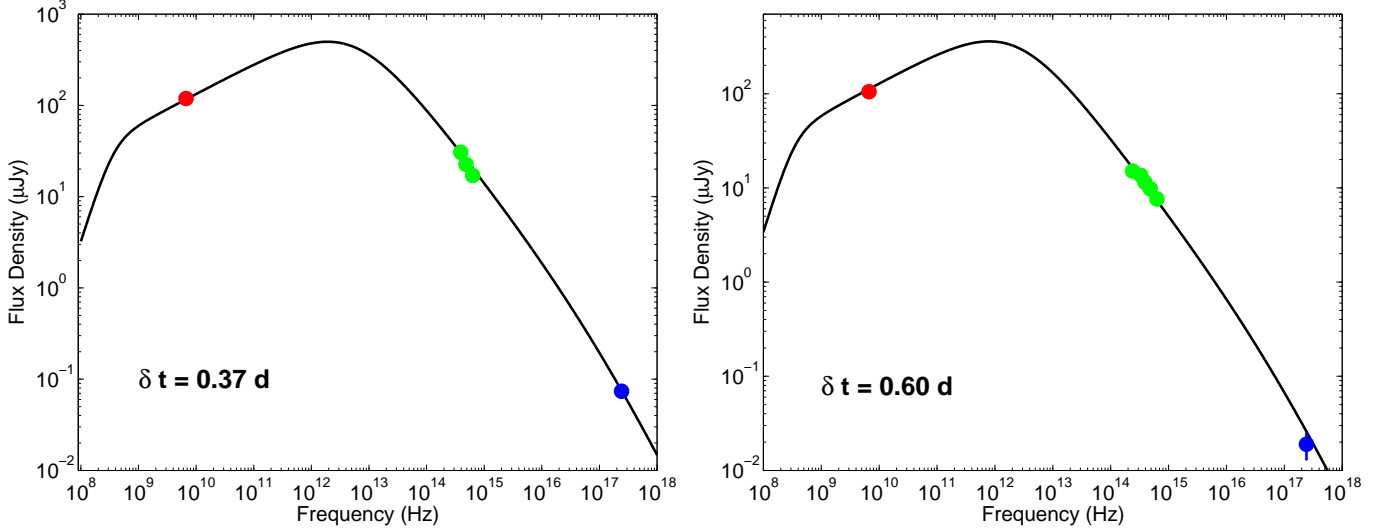
Greisen 2003) for data calibration and analysis. We detect a source located at RA=11<sup>h</sup>28<sup>m</sup>48.15<sup>s</sup>, Dec=+17°04′18.0″ (J2000;  $\delta\text{RA}=0.21''$ ,  $\delta\text{Dec}=0.14''$ ), consistent with the optical and X-ray afterglow positions. We obtained three subsequent 5.8 GHz observations at  $\delta t \approx 1.4$  d,  $\approx 4.3$  d, and  $\approx 84.3$  d in which the source faded, indicating that it is the radio afterglow. We also obtained observations at a mean frequency of 21.8 Hz at  $\delta t \approx 1.4$  d (using J1118+1234 as the phase calibrator), where the afterglow is not detected.

We measure flux densities for the upper and lower sidebands from each epoch using AIPS/JMFIT, and calculate  $3\sigma$

upper limits from source-free regions using AIPS/IMSTAT. The radio afterglow detections and upper limits are listed in Table 1.

### 3. BROAD-BAND AFTERGLOW PROPERTIES

We utilize the broad-band afterglow observations to constrain the explosion properties and circumburst environment of GRB 130603B. We emphasize that this is only the third radio afterglow detection of a short GRB, thereby enabling substantially tighter constraints on the physical properties than for the majority of short GRBs. We adopt the stan-



**Figure 1.** Radio (red; 6.7 GHz), optical/NIR (green; *grizJ*-band; Berger et al. 2013; Cucchiara et al. 2013; Tanvir et al. 2013) and X-ray (blue; 1 keV) afterglow spectral energy distributions of GRB 130603B at  $\delta t = 0.37$  d (left) and 0.60 d (right). The *grizJ* afterglow observations are corrected for  $A_V^{\text{host}} = 1$  mag. Error bars correspond to  $1\sigma$  confidence. A representative best-fit model (black line) is shown in each panel.

standard synchrotron model for a relativistic blastwave in a constant density medium (ISM; Sari et al. 1999; Granot & Sari 2002). This model provides a mapping from the broad-band afterglow flux densities to physical parameters of the explosion and circumburst environment: isotropic-equivalent kinetic energy ( $E_{K,\text{iso}}$ ), circumburst density ( $n_0$ ), fractions of post-shock energy in radiating electrons ( $\epsilon_e$ ) and magnetic fields ( $\epsilon_B$ ), and the electron power-law distribution index ( $p$ ), with  $N(\gamma) \propto \gamma^{-p}$  for  $\gamma \gtrsim \gamma_{\text{min}}$ .

To determine the locations of the break frequencies with respect to the X-ray, optical and radio bands, we first compare the spectral indices from the afterglow observations. From the X-ray spectral fit, we find  $\beta_X \equiv 1 - \Gamma = -1.2 \pm 0.1$  (Section 2.1), while from the optical *grizJ*-band afterglow photometry at  $\delta t \approx 0.6$  d (Table 1; Cucchiara et al. 2013; Tanvir et al. 2013), we measure an observed spectral index of  $\beta_{\text{opt,obs}} = -2.0 \pm 0.1$ . The NIR spectral index measured from the *JK*-bands at the same epoch is substantially shallower,  $\beta_{\text{NIR}} = -0.6 \pm 0.2$ , indicating that the optical flux is suppressed by extinction in the host galaxy. To determine the amount of extinction, we apply a Milky Way extinction curve (Cardelli et al. 1989) to the *grizJK* photometry, fitting for the spectral index and rest-frame extinction ( $A_V^{\text{host}}$ ), and find best-fit values of  $\beta_{\text{opt}} = -0.84 \pm 0.10$  and  $A_V^{\text{host}} = 1.0 \pm 0.1$  mag, similar to the values found from other analyses (de Ugarte Postigo et al. 2013; Jin et al. 2013). Using standard Galactic relations between the intrinsic hydrogen column density and rest-frame extinction (Predehl & Schmitt 1995; Watson 2011), we find an inferred extinction of  $A_V^{\text{host}} = 0.9 - 1.6$  mag from the  $N_{\text{H,int}}$  value determined in Section 2.1, consistent with  $A_V^{\text{host}} \approx 1$  mag. A comparison of  $\beta_{\text{opt}}$  and  $\beta_X$  indicates that the cooling frequency,  $\nu_c$ , lies between the optical and X-ray bands, and that  $p = 2.55 \pm 0.15$ .

To determine the location of the self-absorption and peak frequencies ( $\nu_a$  and  $\nu_m$ , respectively) with respect to the optical and radio bands, we compare the radio spectral slope determined from VLA observations at  $\delta t \approx 0.37$  d and 1.43 d (Table 1) to  $\beta_{\text{opt}}$ . We find that  $\beta_{\text{rad}}$  does not match the optical slope and furthermore is consistent with  $F_\nu \propto \nu^{1/3}$ , as expected for  $\nu_a < \nu_{\text{rad}} < \nu_m$ .

Knowing the relative locations of the break frequencies with respect to our observing bands, we use the standard synchrotron model to determine allowable ranges for the physical parameters. We use the afterglow SED at two common epochs,  $\delta t = 0.37$  d and 0.60 d, where the optical/NIR fluxes are corrected for  $A_V^{\text{host}} = 1$  mag (Figure 1). We also use the constraints  $\epsilon_e, \epsilon_B < 1/3$  and determine:

$$3.3 \times 10^8 < \nu_a < 2.6 \times 10^9 \text{ Hz} \quad (1)$$

$$5.8 \times 10^{-2} < \epsilon_e < 1/3 \quad (2)$$

$$2.0 \times 10^{-3} < \epsilon_B < 1/3 \quad (3)$$

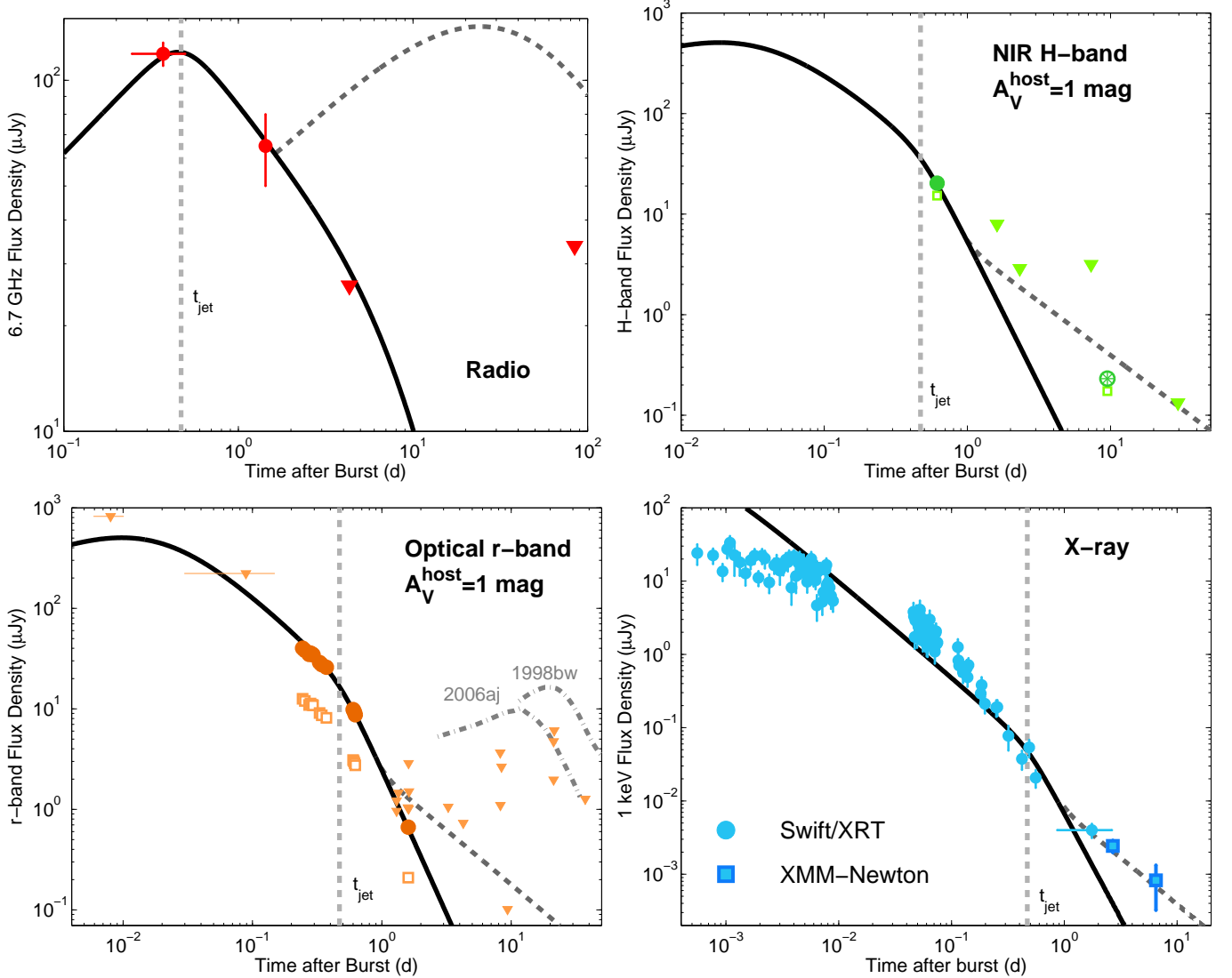
$$0.6 \times 10^{51} \text{ erg} < E_{K,\text{iso}} < 1.7 \times 10^{51} \text{ erg} \quad (4)$$

$$4.9 \times 10^{-3} \text{ cm}^{-3} < n_0 < 30 \text{ cm}^{-3} \quad (5)$$

where the ranges are set by the uncertainty in  $\nu_a$  and the conditions that  $\epsilon_e, \epsilon_B < 1/3$ . We note that the inferred isotropic-equivalent kinetic energy is comparable to  $E_{\gamma,\text{iso}} \approx 2.1 \times 10^{51}$  erg. We show representative model SEDs along with the broad-band afterglow observations in Figure 1.

#### 4. AFTERGLOW EVOLUTION: JET BREAK

To investigate the temporal behavior of the optical/NIR afterglow, we interpolate all of the available afterglow data (Table 1) to the optical *r*-band using the observed spectral slope (Section 3) and then correct these fluxes for  $A_V^{\text{host}} = 1$  mag. The resulting temporal behavior of the optical afterglow flux is characterized by a broken power law (Figure 2). Invoking a broken power law model with two segments, we find pre- and post-break decay indices of  $\alpha_1 \approx -1.26 \pm 0.05$  and  $\alpha_2 = -2.73 \pm 0.08$ , with a break time of  $t_b = 0.47_{-0.06}^{+0.02}$  d. The required change in the temporal index is therefore  $\Delta\alpha \approx 1.5$ . Although there are several possibilities that can explain breaks in GRB afterglow light curves, most of these cannot explain the large change in slope and the steep post-break decline seen here. For instance, the transition of the cooling frequency across the band predicts  $\Delta\alpha = 0.25$  (Sari et al. 1998), while the cessation of energy injection observed in long GRBs typically leads to  $\Delta\alpha \approx 0.7$  with moderate decline rates of



**Figure 2.** Radio through X-ray afterglow light curves of GRB 130603B. Error bars correspond to  $1\sigma$  confidence, and triangles denote  $3\sigma$  upper limits. The afterglow model is shown as a black line, while the jet break time of  $t_j \approx 0.47$  d is marked by a vertical grey dashed line. Also shown is a model with energy injection (dark grey dashed line) that fits the X-ray excess but violates the detections and limits in the optical, NIR, and radio bands. *Top left:* 6.7 GHz observations with the VLA (red). *Top right:* H-band observations (green; Berger et al. 2013; de Ugarte Postigo et al. 2013; Tanvir et al. 2013), where JK-band observations are extrapolated to H-band using  $\beta_{\text{opt}} = -2$ . The observed values (open green squares) are corrected for  $A_V^{\text{host}} = 1$  mag (filled green symbols). The circled asterisk at  $\delta t \approx 9$  d is the “kilonova” associated with GRB 130603B (Berger et al. 2013; Tanvir et al. 2013). *Bottom left:* Optical r-band observations (orange; this work, Cucchiara et al. 2013; de Ugarte Postigo et al. 2013; Tanvir et al. 2013), where *giz*-band observations are extrapolated to r-band using  $\beta_{\text{opt}} = -2$ . The observed values (open orange squares) are corrected for  $A_V^{\text{host}} = 1$  mag (filled orange symbols). The displayed upper limits (orange triangles) are also corrected for extinction. Also shown are the optical light curves of GRB-SN 1998bw (Galama et al. 1998; Clocchiatti et al. 2011) and GRB-SN 2006aj (dot-dashed lines; Mirabal et al. 2006) corrected for extinction and redshifted to  $z = 0.3565$ . *Bottom right:* Observations from *Swift*/XRT (blue circles) and *XMM-Newton* (blue squares) at 1 keV.

$\alpha_1 \approx -0.5$  and  $\alpha_2 \approx -1.2$  (Nousek et al. 2006; Zhang et al. 2006; Liang et al. 2007). A steep drop in the density is predicted to cause maximum changes of  $\Delta\alpha \approx 0.4$  for density contrasts of  $\sim 10$  (Nakar & Granot 2007), and would require a density contrast of  $\gtrsim 1000$  to account for  $\Delta\alpha = 1.5$ .

We therefore conclude that the temporal steepening is instead a jet break, when the edge of a relativistically-beamed outflow becomes visible to the observer (Sari et al. 1999; Rhoads 1999). In this scenario, the post-break flux declines as  $t^{-p}$  (Sari et al. 1999). Indeed, we find good agreement between  $\alpha_2 = -2.7 \pm 0.1$  determined from the optical light curve and  $p = 2.55 \pm 0.15$  independently determined from the broadband SED (Section 3).

In addition, since the radio band lies between  $\nu_a$  and  $\nu_m$

(Section 3), the radio flux density should increase as  $F_\nu \propto t^{1/2}$  in the spherical regime (Granot & Sari 2002), while the flux will decrease as  $F_\nu \propto t^{-1/3}$  in a post-jet break scenario (Sari et al. 1999). We find that the observed radio flux of GRB 130603B declines with  $\alpha \approx -0.45$  between  $\delta t \approx 0.4$  d and 1.3 d, demonstrating that the evolution is not isotropic (Figure 2). Thus, the temporal behavior of both the optical and radio afterglows support a jet break at  $t_j \approx 0.47$  d, making this the first detection of a jet break in the radio afterglow of a short GRB.

In conjunction with the energy, density, and redshift, the time of the break can be converted to a jet opening angle (Sari et al. 1999; Frail et al. 2001),

$$\theta_j = 9.51 t_{j,d}^{3/8} (1+z)^{-3/8} E_{K,iso,52}^{-1/8} n_0^{1/8} \text{ deg} \quad (6)$$

where  $t_{j,d}$  is in days,  $E_{K,iso,52}$  is in units of  $10^{52}$  erg and  $n_0$  is in units of  $\text{cm}^{-3}$ . For the ranges of  $E_{K,iso}$  and  $n_0$  in Equations 4-5, we calculate  $\theta_j \approx 4 - 14^\circ$ . However, given a NS-NS/NS-BH progenitor as indicated by the likely detection of a kilonova (Berger et al. 2013; Tanvir et al. 2013), the circumburst density is likely more typical of ISM densities ( $n \lesssim 1 \text{ cm}^{-3}$ ), leading to  $\theta_j \approx 4 - 8^\circ$ .

### 5. LATE-TIME X-RAY EXCESS

The synchrotron afterglow model with a jet break at  $t_j \approx 0.47$  d provides a good match to the radio and optical light curves (Figure 2). However, unlike the optical afterglow behavior, we do not observe significant steepening in the X-ray light curve, and instead the afterglow flux at  $\delta t \gtrsim 0.03$  d can be characterized by a single power law with  $\alpha_X = -1.88 \pm 0.15$ . Thus, our afterglow model underpredicts the X-ray flux by a factor of  $\approx 5$  at  $\delta t \gtrsim 2$  d (see Figure 2). We note that Tanvir et al. (2013) and de Ugarte Postigo et al. (2013) previously claimed that the same broken power-law fits both the optical and X-ray data, but this was based on the XRT data alone. The *XMM-Newton* observations do not support these claims. We are thus motivated to consider an additional energy source for this excess X-ray emission, which follows  $L_X \simeq 4 \times 10^{43} t_d^{-1.88} \text{ erg s}^{-1}$  (where  $t_d$  is in units of days).

One possible source of late-time excess X-ray emission is continued energy injection into the blastwave from ongoing central engine activity (Zhang & Mészáros 2002; Zhang et al. 2006). We first consider continuous power law energy injection with  $L \propto t^q$ . We find that energy injection beginning at  $\delta t \approx 1$  d characterized by  $q = 0.3$  (e.g.,  $E \propto t^{1.3}$ ) provides an adequate fit to the late-time X-ray light curve (Figure 2). Assuming that the injection extends to  $\delta t \approx 10$  d, the energy injection factor is  $\approx 9.5$ . However, this energy injection model violates both the optical and radio light curves; in particular, the radio upper limits are a factor of  $\approx 3$  below the predicted flux with energy injection (Figure 2). Therefore, continuous energy injection is not a viable explanation for the excess X-ray emission.

We are therefore motivated to consider a source of emission that predominantly contributes in the X-rays with negligible effects on the other bands. We focus on scenarios in which this emission originates from the central engine, but which we now assume can be viewed directly through the merger ejecta (we justify this assumption below). We first consider ongoing accretion (“fall-back”) onto the newly-formed black hole following a compact object merger. By extrapolating the matter trajectories from numerical simulations of the merger process to late times, Rosswog (2007) predicts accretion luminosities of  $L_{acc} \sim 10^{43} - 10^{45} \text{ erg s}^{-1}$  on a timescale of  $\sim 1$  d, with a temporal decay of  $L_{acc} \sim t^{-\alpha}$  similar to the canonical prediction of  $\alpha = -5/3$  for a tidally disrupted star (Rees 1988). Assuming a radiative efficiency of  $\gtrsim 10\%$ , this scenario is well matched to the X-ray light curve of GRB 130603B. However, such efficiencies are optimistic and more detailed models of the fall-back process from compact object mergers (Rossi & Begelman 2009) predict much lower X-ray luminosities of  $L_X \lesssim 10^{-3} L_{acc}$ , which would not be large enough to explain the observed excess.

Another possibility is that the X-ray emission is powered by the spin-down of a massive magnetar remnant (Metzger et al.

2008; Bucciantini et al. 2012; Zhang 2013), a process that has been used to explain putative plateaus in the X-ray afterglows of short GRBs (Rowlinson et al. 2013). Such remnants may be at least temporarily stable to gravitational collapse, if they rotate with spin periods of  $\lesssim$  few ms (e.g. Özel et al. 2010; Giacomazzo & Perna 2013). The remnant may also acquire strong magnetic fields of  $\gtrsim 10^{14} - 10^{15}$  G, similar to those of Galactic magnetars (Duncan & Thompson 1992; Zrake & MacFadyen 2013). The magnetar model predicts that the late-time spin-down luminosity should decay as  $L_{sd} \propto t^{-2}$ , consistent with the observed temporal decay of  $\alpha_X = -1.88 \pm 0.15$ . The predicted spectrum is  $F_\nu \propto \nu^{-1}$  (Metzger et al. 2013), also consistent with the observed spectral index  $\beta_X \approx -1.2 \pm 0.1$ . Fitting the entire X-ray light curve with a magnetar model characterized by the duration and luminosity of the plateau (Zhang & Mészáros 2002), and assuming  $M_{NS} = 1.4 - 2.5 M_\odot$  and  $R_{NS} = 10^6$  cm, we find best-fit parameters of  $B \approx 2 \times 10^{16}$  G and  $P \approx 15 - 25$  ms (where higher mass corresponds to slower spin periods). However, such slow spin periods are likely unphysical in the merger context due to the substantial angular momentum of the initial binary. Instead, assuming a more reasonable initial spin period of 1 ms, the required magnetic field strength to produce the observed X-ray luminosity at  $\gtrsim 1$  d assuming 10% radiative efficiency, is  $B \approx 10^{15}$  G, but such a model would under-predict the light curve at  $\delta t \lesssim 3000$  s by a factor of a few. We thus conclude that the magnetar scenario could potentially explain the late-time X-ray excess for  $\delta t \gtrsim 3000$  s.

To justify that we can observe the central engine directly, as is required in either the fall-back or magnetar models, the merger ejecta must be transparent to soft X-rays. Due to the high bound-free X-ray opacity of neutral matter, this in turn requires that the engine be sufficiently luminous to reionize the merger ejecta (Metzger et al. 2013). The ejecta becomes transparent to X-rays once two conditions are satisfied: (1) the opacity becomes dominated by electron scattering ( $\kappa_{bf}/\kappa_{es} \lesssim 1$ , where  $\kappa_{bf}$  and  $\kappa_{es}$  are the bound-free and electron scattering opacities, respectively) and (2) the electron scattering optical depth,  $\tau_{es}$ , through the ejecta decreases to  $\lesssim 1$ . Assuming ejecta mainly composed of hydrogen-like iron<sup>9</sup> and an ejecta temperature of  $T_{ej} = 10^4$  K, we derive the following expressions (see Appendix for details):

$$\frac{\kappa_{bf}}{\kappa_{es}} \approx 0.13 \left( \frac{L_X}{4 \times 10^{43} \text{ erg s}^{-1}} \right)^{-1} \left( \frac{M_{ej}}{10^{-3} M_\odot} \right) \left( \frac{t}{1 \text{ d}} \right)^{-1} \left( \frac{v_{ej}}{c} \right)^{-1}, \quad (7)$$

$$\tau_{es} = \rho_{ej} \kappa_{es} R_{ej} \simeq 0.02 \left( \frac{M_{ej}}{10^{-3} M_\odot} \right) \left( \frac{v_{ej}}{c} \right)^{-2} \left( \frac{t}{1 \text{ d}} \right)^{-2} \quad (8)$$

where  $M_{ej}, R_{ej} = v_{ej} t$ , and  $v_{ej}$  are the (effective) mass, radius, and velocity of the ejecta along the observer line of sight, respectively. Using fiducial values for  $M_{ej}$  and  $v_{ej}$  of the merger (Hotokezaka et al. 2013; Fernández & Metzger 2013), we find that both conditions are satisfied on timescales  $\delta t \gtrsim$  few hr given the observed X-ray luminosity. We thus conclude that the ejecta are indeed transparent to the soft X-rays at late times, supporting the idea that direct radiation from the central engine could produce the observed X-ray excess emission

<sup>9</sup> Outflows along the polar direction arise chiefly from the accretion disk and are expected to be composed of Fe-like nuclei (e.g. Metzger et al. 2008).



in GRB 130603B.

We note that of the  $\approx 10$  short GRBs with X-ray observations to  $\delta t \gtrsim 1$  d, two events, GRBs 050724 and 080503, also exhibited late-time X-ray excess emission on timescales of  $\delta t \sim 1-2$  d (Grupe et al. 2006; Perley et al. 2009). However, unlike GRB 130603B, these bursts both had corresponding behavior in the optical bands (Malesani et al. 2007; Perley et al. 2009), suggesting that the optical and X-ray emission were from the same emitting region.

## 6. COMPARISON TO PREVIOUS SHORT GRBS

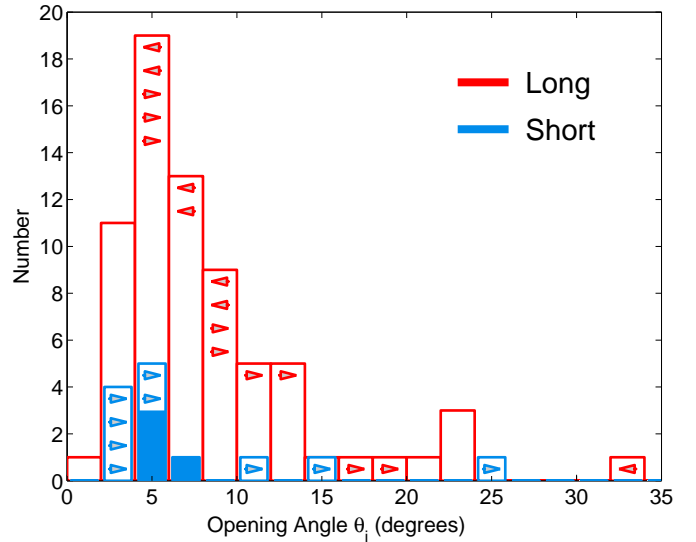
The broad-band afterglow observations of GRB 130603B provide the second detection of a multi-wavelength jet break in a short GRB, and the first detection of a jet break in the radio band. Radio afterglow emission has thus far been detected in two short GRBs: GRB 050724A (Berger et al. 2005), GRB 051221A (Soderberg et al. 2006). The ability to monitor the radio afterglow of GRB 130603B at a flux density level of  $\lesssim 0.1$  mJy highlights the improved sensitivity of the VLA.

Indeed, the radio evolution can provide an important constraint on the progenitor. In the context of the compact object binary progenitor, the radioactive decay of  $r$ -process elements in the sub-relativistic merger ejecta is predicted to produce transient emission, termed a “kilonova” (Li & Paczyński 1998; Metzger et al. 2010; Goriely et al. 2011; Roberts et al. 2011; Rosswog et al. 2013), which is expected to peak in the NIR (Barnes & Kasen 2013; Kasen et al. 2013; Tanaka & Hotokezaka 2013). Late-time NIR emission in GRB 130603B detected with the *Hubble Space Telescope* (Figure 2; Berger et al. 2013; Tanvir et al. 2013) has been interpreted as the first detection of a  $r$ -process kilonova. An alternative scenario to explain the excess NIR emission of GRB 130603B may be a wide, mildly relativistic component of a structured jet (Jin et al. 2013) which has been used to explain the light curve behavior of a handful of long GRBs (e.g., Berger et al. 2003; Sheth et al. 2003; Peng et al. 2005; Racusin et al. 2008). In this scenario, the predicted radio emission is similarly boosted, and will be  $\approx 80 \mu\text{Jy}$  at  $\delta t \approx 84$  d, the time of our final radio observations (Jin et al. 2013). Instead, the non-detection of any radio emission to  $\lesssim 34 \mu\text{Jy}$  provides a strong constraint on the existence of a two component jet, and supports the kilonova interpretation of the NIR emission.

The detection of a jet break in GRB 130603B leads to an opening angle measurement of  $4-14^\circ$ , with a more likely range of  $4-8^\circ$ . This opening angle is the fourth<sup>10</sup> such measurement for a short GRB after GRB 051221A ( $7^\circ$ ; Soderberg et al. 2006), GRB 090426 ( $5-7^\circ$ ; Nicuesa Guelbenzu et al. 2011), and GRB 111020A ( $3-8^\circ$ ; Fong et al. 2012). From these four short GRB opening angle measurements, the median is  $\langle \theta_j \rangle \approx 6^\circ$  (Figure 3).

The non-detections of jet breaks can provide lower limits on the opening angles, assuming on-axis orientation, as off-axis observing angles could disguise jet breaks (van Eerten & MacFadyen 2012, 2013). Indeed, such non-detections to timescales of  $\sim 1$  day with *Swift*/XRT have led to lower limits of  $\theta_j \gtrsim 2-6^\circ$  (Fong et al. 2012), while monitoring with more sensitive instruments such as *Chandra* and *XMM-Newton* to timescales of  $\sim 1$  week has led to more meaningful limits of  $\theta_j \gtrsim 10-25^\circ$  (Figure 3; Fox et al.

<sup>11</sup> We note that Nicuesa Guelbenzu et al. (2012) claimed a jet break in the GRB 090305A afterglow but this is based on a single optical data point.



**Figure 3.** Distribution of opening angles for long (red) and short (blue) GRBs, updated from Fong et al. (2012). Arrows represent upper and lower limits. The long GRB population includes pre-*Swift* (Frail et al. 2001; Berger et al. 2003; Bloom et al. 2003; Ghirlanda et al. 2004; Friedman & Bloom 2005), *Swift* (Racusin et al. 2009; Filgas et al. 2011), and *Fermi* (Cenko et al. 2010; Goldstein et al. 2011; Cenko et al. 2011) bursts. For short GRBs, the existing measurements are GRB 051221A ( $7^\circ$ ; Soderberg et al. 2006), GRB 090426 ( $5-7^\circ$ , assigned  $6^\circ$  here; Nicuesa Guelbenzu et al. 2011), GRB 111020A ( $3-8^\circ$ , assigned  $5.5^\circ$  here; Fong et al. 2012) and GRB 130603B ( $4-8^\circ$ , assigned  $6^\circ$  here; this work). Short GRB lower limits are from the non-detection of jet breaks in *Swift*/XRT data (Fong et al. 2012), *Chandra* data for GRBs 050724A (Grupe et al. 2006), 101219A (Fong et al. 2013), 111117A (Margutti et al. 2012a; Sakamoto et al. 2013), and 120804A (Berger et al. 2013) and optical data for GRBs 050709 (Fox et al. 2005) and 081226A (Nicuesa Guelbenzu et al. 2012).

2005; Grupe et al. 2006; Berger et al. 2013). The search for jet breaks has been less fruitful in the optical bands, primarily due to the intrinsic faintness of the optical afterglows and contamination from host galaxies. Indeed, the sole lower limit from a well-sampled optical light curve is from GRB 081226A, with  $\theta_j \gtrsim 3^\circ$  (Nicuesa Guelbenzu et al. 2012), while we conservatively adopt a lower limit of  $\theta_j \gtrsim 15^\circ$  for GRB 050709 based on a sparsely-sampled optical light curve (Fox et al. 2005). Using the measured opening angles and lower limits of  $\gtrsim 10-25^\circ$ , a likely median for short GRBs is  $\langle \theta_j \rangle \approx 10^\circ$ .

The opening angle distribution of short GRBs impacts the true energy scale and event rate. The true energy is lower than the isotropic-equivalent value by the beaming factor,  $f_b$  ( $f_b \equiv 1 - \cos(\theta_j)$ ,  $E = f_b E_{\text{iso}}$ ), while the actual event rate is increased by  $f_b^{-1}$ . For GRB 130603B, with an opening angle of  $\approx 4-8^\circ$ , the resulting beaming factor is  $f_b \approx (0.2-1) \times 10^{-2}$ . Therefore, the true energies are  $E_\gamma \approx (0.5-2) \times 10^{49}$  erg and  $E_K \approx (0.1-1.6) \times 10^{49}$  erg. The small population of short GRBs with well-constrained opening angles have beaming-corrected energies of  $E_\gamma \approx E_K \approx 10^{49}$  erg (Soderberg et al. 2006; Burrows et al. 2006; Grupe et al. 2006; Fong et al. 2012; Nicuesa Guelbenzu et al. 2012), roughly two orders of magnitude below the inferred true energies for long GRBs (Frail et al. 2001; Bloom et al. 2003; Kocevski & Butler 2008; Racusin et al. 2009).

The true event rate is elevated compared to the observed rate by  $f_b^{-1}$ . The current estimated observed short GRB volumetric rate is  $\approx 10 \text{ Gpc}^{-3} \text{ yr}^{-1}$  (Nakar et al. 2006). For a median opening angle of  $\approx 10^\circ$ , the median in-



verse beaming factor is  $f_b^{-1} \approx 70$ , resulting in a true rate of  $\approx 700 \text{ Gpc}^{-3} \text{ yr}^{-1}$ . The observed all-sky event rate of  $\approx 0.3 \text{ yr}^{-1}$  within 200 Mpc (Guetta & Piran 2005) then becomes  $\approx 20 \text{ yr}^{-1}$ . This rate is comparable to estimates for NS-NS merger detections with Advanced LIGO/VIRGO (LIGO Scientific Collaboration et al. 2013).

## 7. CONCLUSIONS

We presented broad-band observations of the afterglow of GRB 130603B, uncovering a jet break in the optical and radio light curves at  $\delta t \approx 0.47 \text{ d}$ . This comprehensive data set marks the first detection of a jet break in the radio band and the third radio afterglow detection in nearly a decade of follow-up. The inferred opening angle is  $\theta_j \approx 4-8^\circ$ , leading to true energy releases of  $E_\gamma \approx E_K \approx 10^{49} \text{ erg}$ .

We observe excess X-ray emission at  $\gtrsim 1 \text{ day}$  with no corresponding emission in the other bands. We rule out energy injection from ongoing activity from the central engine due to the non-detection of any radio or optical emission on a similar timescale. We find that fall-back accretion can explain the late-time excess only if the radiative efficiency is  $\gtrsim 10\%$ . Finally, we consider that the emission is due to the spin-down of a massive magnetar and find that a model characterized by a spin period of  $\approx 1 \text{ ms}$  and magnetic field of  $\approx 10^{15} \text{ G}$  provides a good fit to the emission at  $\delta t \gtrsim 1 \text{ d}$ , but underpredicts the X-ray emission at  $\lesssim 3000 \text{ s}$ . Furthermore, we show that

the merger ejecta are transparent to soft X-rays (also see Appendix), ensuring that the engine can be viewed in X-rays.

GRB 130603B highlights the importance of multi-wavelength afterglow observations, which provide the only route to constraints on the basic explosion properties of GRBs. In particular, the radio band is advantageous because unlike the optical, it does not typically suffer from host galaxy contamination, and can provide an additional constraint on the circumburst density. Thus, continued target-of-opportunity efforts at the VLA will provide invaluable information on the sub-parsec explosion environments. In addition, the non-detection of late-time radio emission, coupled with the detection of NIR excess emission, can provide unambiguous support for the kilonova interpretation for future events, as it has for GRB 130603B.

The opening angle determination for GRB 130603B is the fourth robust jet break measurement for a short GRB. Using realistic assumptions for the opening angle distribution, this implies a conservative volumetric event rate of  $\approx 700 \text{ Gpc}^{-3} \text{ yr}^{-1}$ , and an all-sky event rate of  $\approx 20 \text{ yr}^{-1}$  within 200 Mpc, consistent with the predictions of NS-NS merger detections with Advanced LIGO/VIRGO. However, the opening angle distribution for wider jets of  $\gtrsim 5^\circ$  is poorly constrained, and it is necessary to continue monitoring short GRB afterglows to late times to characterize this part of the distribution.

## APPENDIX

### RE-IONIZATION MODEL

The average density of the freely-expanding ejecta decreases with time as  $\rho_{\text{ej}} \simeq M_{\text{ej}}/(4\pi/3R_{\text{ej}}^3)$ , where  $M_{\text{ej}}, R_{\text{ej}} = v_{\text{ej}}t$  is the ejecta radius, and  $v_{\text{ej}}$  is the ejecta velocity. The ionization state of the ejecta is determined by comparing the absorption rate of ionizing photons  $\mathcal{R}_{\text{ion}} = C n_{\gamma, \nu > \nu_1} \sigma_{\nu_1} c$  (per ion) to the rate of recombination  $\mathcal{R}_{\text{rec}} = n_e \alpha_{\text{rec}}$ , where  $n_{\gamma, \nu > \nu_1} = L_X/4\pi h\nu_1 R_{\text{ej}}^2 c$  is the number density of ionizing photons;  $C$  is a constant of order unity that depends on the spectrum of the ionizing radiation;  $L_{\nu_1}$  is the specific X-ray luminosity near the ionization threshold energy  $h\nu_1 \sim 10 \text{ keV}$ ;  $\sigma_{\nu_1} \simeq 8 \times 10^{-21} \text{ cm}^2$  is the photoionization cross section at  $\nu = \nu_1$  and  $n_e \simeq \rho_{\text{ej}}/2m_p$  is the number density of electrons in the ejecta;  $\alpha_{\text{rec}} \approx 2.0 \times 10^{-10} T_4^{-0.8} \text{ cm}^3 \text{ s}^{-1}$  is the [type 2] recombination coefficient (e.g. Osterbrock & Ferland 2006); and  $T_4 \sim 1$  is the ejecta temperature in units of  $10^4 \text{ K}$ . The above results can be combined to determine the ratio of the bound-free  $\kappa_{\text{bf}} = f_n \sigma_{\nu_1}/26m_p$  and electron scattering  $\kappa_{\text{es}} = 0.2 \text{ cm}^2 \text{ g}^{-1}$  opacities of the ejecta, where  $f_n \approx \mathcal{R}_{\text{rec}}/\mathcal{R}_{\text{ion}} \ll 1$  is the neutral fraction (set by the balance between ionization and recombination rates). The resulting coefficient is in Equation 7, from which it is shown that the ejecta is sufficiently ionized for the observed X-rays to originate directly from the central engine interior to the ejecta.

One might be concerned that an X-ray source of sufficient luminosity to ionize the ejecta along the observer line of sight (perpendicular to the merger plane) would also be sufficient to ionize matter ejected in the equator, the radioactive heating of which powers the optical/NIR kilonova emission. Indeed, the red colors of the kilonova result from the high line opacity of the lanthanides, which would vanish were the ejecta ionized by the central engine. However, equation (7) shows that  $\kappa_{\text{bf}}/\kappa_{\text{es}} \propto L_X^{-1} t^{-1}$  is larger for the equatorial ejecta (due to its larger mass  $M_{\text{ej}}$  and lower velocity  $v_{\text{ej}}$ ), and that this ratio increases with time  $\propto t$  (since  $L_X \propto t^{-2}$ , approximately). The fact that  $\kappa_{\text{bf}}/\kappa_{\text{es}}$  may be significantly  $\gg 1$  for the equatorial kilonova ejecta implies that the latter may remain neutral, preserving the kilonova emission. However, more detailed calculations, including the different (and uncertain) recombination rates of the lanthanides, is necessary to verify this conclusion.

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